

GRB060218 AS A TIDAL DISRUPTION OF A WHITE DWARF BY AN INTERMEDIATE MASS BLACK HOLE

ROMAN V. SHCHERBAKOV^{1,2,3}, ASAF PE'ER^{4,5}, CHRISTOPHER S. REYNOLDS^{1,2}, ROLAND HAAS^{6,7}, TANJA BODE⁷, PABLO LAGUNA⁷

Draft version December 17, 2012

ABSTRACT

A highly unusual pair of a gamma-ray burst (GRB) GRB060218 and an associated supernova SN2006aj has puzzled theorists for years. A supernova shock breakout and a jet from a newborn stellar mass compact object were put forward to explain its multiwavelength signature. We propose that the source is naturally explained by another channel, a tidal disruption of a white dwarf (WD) by an intermediate mass black hole (IMBH). The tidal disruption is accompanied by a tidal pinching, which leads to the ignition of a WD and a supernova. Some debris falls back onto the IMBH, forms a disk, which quickly amplifies the magnetic field, and launches a jet. We successfully fit soft X-ray spectrum with the Comptonized blackbody emission from a jet photosphere. The optical/UV emission is consistent with self-absorbed synchrotron from the expanding jet front. The accretion rate temporal dependence $\dot{M}(t)$ in a tidal disruption provides a good fit to soft X-ray lightcurve. The IMBH mass is found to be about $10^4 M_\odot$ in three independent estimates: (1) fitting tidal disruption $\dot{M}(t)$ to soft X-ray lightcurve; (2) computing the jet base radius in a jet photospheric emission model; (3) inferring the central BH mass based on a host dwarf galaxy stellar mass. The supernova position is consistent with the center of the host galaxy, while low supernova ejecta mass is consistent with a WD mass. High expected rate of tidal disruptions in dwarf galaxies is consistent with one source observed by *Swift* satellite over several years at GRB060218 distance of 150 Mpc. The encounters with the WDs provide a lot of fuel for IMBH growth.

Subject headings: accretion – black hole physics – gamma rays: bursts – radiation mechanisms: general – supernovae: general – X-rays: individual (GRB060218)

1. INTRODUCTION

The existence of the stellar mass black holes (BHs) with mass $M_{BH} < 100M_\odot$ and the supermassive black holes (SMBH) with mass $M_{BH} > 10^5 M_\odot$ has long been established. A population of intermediate mass black holes (IMBH) likely exists. Those BHs have masses in a range $100M_\odot < M_{BH} < 10^5 M_\odot$. They could live in the centers of dwarf galaxies (Dong et al. 2007; Greene 2012) or globular clusters (GCs) (Fabbiano et al. 1997, 2001; Colbert & Mushotzky 1999; Matsumoto et al. 2001; Gültekin et al. 2004). The IMBHs can form by a collapse of a massive cloud (Begelman et al. 2006) or a massive star (Fryer et al. 2001; Madau & Rees 2001; Schneider et al. 2002) or grow from a stellar mass BH. A star cluster G1 in M31 galaxy is estimated to host a $2 \times 10^4 M_\odot$ BH (Gebhardt et al. 2002) based on the velocity dispersion profile. SDSS J160531.84+174826.1 dwarf galaxy is estimated to have an IMBH with a mass $\sim 7 \times 10^4 M_\odot$ based on luminosity scaling relations (Dong et al. 2007). One of the best candidates is HLX-1 source in ESO 243-49 galaxy. Based on a thin disk thermal state the estimated BH mass is $3 \times 10^4 M_\odot$ (Davis et al. 2011). All those candidates are still tentative and the uncertainty of the mass estimates is up to an order of magnitude. More IMBH candidates with qualitatively differ-

ent observational signatures may provide stronger evidence for existence of such objects.

Tidal disruptions of the white dwarfs (WDs) by IMBHs provide such qualitatively different signature. Unlike disruptions of main sequence (MS) stars, which happen far away from the BHs and have slow timescales, the disruptions of WDs are very fast. Such disruptions may lead to accretion rates up to $10^4 M_\odot \text{yr}^{-1}$ (Haas et al. 2012). A flow with an extreme accretion rate may produce a short powerful burst of radiation, for example, when a jet is launched. The X-ray emission of *Swift* J1644+57 source is attributed to a jet launched by a super-Eddington accretion disk formed after a tidal disruption (Bloom et al. 2011; Burrows et al. 2011; Levan et al. 2011; Zauderer et al. 2011).

When particle acceleration is inefficient, then the only radiation from the jet is the photospheric emission (Eichler & Levinson 2000; Mészáros & Rees 2000; Daigne & Mochkovitch 2002; Rees & Mészáros 2005). The radiation field is in a thermal equilibrium with matter within a dense hot jet up to the distance R_{ph} along the jet, called a photospheric radius, where the optical depth to Compton scattering is about unity $\tau_\sigma \sim 1$. The photospheric emission has a quasi-blackbody spectrum (Pe'er et al. 2007; Pe'er 2008; Pe'er & Ryde 2011; Beloborodov 2011). Particle acceleration may happen either in the entire volume of the jet or in the internal shocks. Dissipation in the volume of the jet, in particular, in the sub-photospheric regions leads to a modified blackbody spectrum from near the jet photosphere (Pe'er et al. 2005, 2006; Giannios 2006, 2012). Modifications include the Comptonization and the broadening of the blackbody peak with pair production. Effective particle acceleration in the internal shocks (Meszaros 2006) is likely responsible for the emission of gamma-ray bursts (GRBs). When two shells moving at different speeds collide above the photosphere, a shock forms, and a substantial fraction of rela-

¹ Department of Astronomy, University of Maryland, College Park, MD 20742, USA

² Joint Space Science Institute, University of Maryland, College Park MD 20742, USA

³ Hubble Fellow

⁴ Harvard-Smithsonian Center for Astrophysics, 60 Garden Street, Cambridge, MA 02138, USA

⁵ Department of Physics, University College Cork, Cork, Ireland

⁶ Theoretical Astrophysics Including Relativity, California Institute of Technology, Pasadena, CA 91125, USA

⁷ Center for Relativistic Astrophysics, School of Physics, Georgia Institute of Technology, Atlanta, GA 30332, USA

tive kinetic energy can be transferred into the electrons (Rees & Meszaros 1994; Sari & Piran 1997), which radiate synchrotron emission. The collisions of shells are especially efficient, when the jet bulk Lorentz factor Γ is large. For the small Lorentz factors and the large photospheric radii, such as in the jets from tidal disruptions of the WDs by the IMBHs, the collisions of shells may happen within the photosphere, and the internal shock signatures might be weaker. The fluctuations of the jet Γ diffuse out, while the jet travels within the photosphere, and the fluctuations outside of the photosphere are small. Even when the slow shells collide, the energy release is weak. Therefore, we expect to see the strong blackbody signature of the photospheric emission. While the energy density of the slow tidal disruption jet might not warrant fast pair production, the Comptonization is still expected to modify the spectrum.

A distinct feature of the WD disruptions by IMBHs is a supernova. As a consequence of the tidal compression along the angular momentum axis, the WD may undergo thermonuclear ignition (Luminet 1985; Luminet & Marck 1985; Luminet & Pichon 1989b) and explode. The energy and the composition of the resulting supernova are the functions of the WD mass m_{WD} , the pericenter radius R_p , and the BH mass M_{BH} (Rossowog et al. 2009). If the WD is massive and the disruption is deep, then the explosion could be similar to a supernova Type Ia with a comparable energy release. Less massive WDs may lead to low luminosity explosions with little *Ni* synthesized. There is no explosion at all in some cases. In any case the ejected mass $\lesssim 1M_\odot$ is less than in the core-collapse supernovae.

In sum, we predict a transient similar to a GRB, but softer and longer, accompanied by a supernova with a small ejecta mass. The best candidate we find in *Swift* GRB catalogue is GRB060218 source. This unusual event is an underluminous very long GRB with a duration $t_{90} \approx 2600$ s and with a smooth X-ray lightcurve (Soderberg et al. 2006). It is accompanied by a fast supernova SN2006aj, which was modeled to have a low ejecta mass $M_{ej} \sim (1-2)M_\odot$ (Mazzali et al. 2006b). The X-ray emission has a blackbody component characteristic of the photospheric radiation. The early X-ray radiation is accompanied by the powerful optical/UV emission (Ghisellini et al. 2007a). Two classes of theories were proposed to explain the source: a supernova shock breakout model (Campana et al. 2006; Waxman et al. 2007; Nakar & Sari 2012) and a model with a mildly relativistic jet from a newborn compact object, such as a magnetar (Fan et al. 2006; Toma et al. 2007; Ghisellini et al. 2007b,a). Nevertheless, we think that a tidal disruption of a WD by an IMBH is not only a viable model for the source, but also more naturally explains some features, such as the duration and the soft quasi-thermal spectrum.

The paper is organized as follows. In § 2 we describe in more detail the observations of GRB060218/SN2006aj. We extensively discuss in § 3 the former theoretical explanations of the source: the shock breakout model and the jet launched by the newborn compact object. Then in § 4 we concentrate on modeling within a tidal disruption scenario. We rederive the dynamics of a tidal disruption and discuss the physics of jet launching. We perform time-resolved spectroscopy of the source and successfully model the X-ray emission with a Comptonized blackbody spectrum from the photosphere. The derived jet base radius corresponds to about $10^4 M_\odot$ IMBH. We propose that the origin of the powerful early optical/UV radiation is the front region of propagating jet. We find that

the full X-ray lightcurve can be fitted well by a scaled dependence of the fallback accretion rate on time $\dot{M}(t)$. The fit provides the IMBH mass about $10^4 M_\odot$. We find that the steep decay phase is consistent with the action of absorption alone and that the afterglow can be naturally explained as powered by the central engine operating in a shallow $t^{-4/3}$ regime. The associated supernova is consistent with a WD origin, while its position is consistent with the center of a host galaxy. The mass of the dwarf host galaxy provides the estimate of a central IMBH mass about $10^4 M_\odot$. In § 5 we estimate the event rates. A high rate of tidal disruptions is predicted in the dwarf galaxies, while the disruptions of the WDs constitute a significant percentage of all disruptions. Since the disruptions of the MS stars mostly happen at large distances from the BHs, then such events could be much dimmer, than the disruptions of the WDs. The tidal disruptions of the WDs provide a plenty of material to feed the central IMBH up to a supermassive size.

2. OBSERVATIONS OF GRB060218

GRB060218 triggered the Burst Alert Telescope (BAT) onboard *Swift* mission satellite on 18 Feb 2006 (Campana et al. 2006). Soon after the trigger the X-ray Telescope (XRT) identified a bright source, whose count rate peaked at around 960 s from BAT initial trigger. Then the source gradually decayed over a continuous observation period, which ended at around 2700 s. The event duration $t_{90} = 2600$ s is unusually long among *Swift* GRBs (Campana et al. 2006). The XRT lightcurve was unusually smooth and regular. The rise period is characterized by hard emission with the BAT flux about equal to the XRT flux (Campana et al. 2006; Toma et al. 2007) and with the peak energy $E_p \sim 5$ keV. The isotropic-equivalent luminosity is $L_{iso} \sim 10^{47} \text{erg s}^{-1}$, which is about 10^5 times less than for the typical GRBs (Toma et al. 2007). The source is found to have a blackbody component with a low temperature $T \sim 0.2$ keV, whose flux contribution increased with time from 15% till 80% (Campana et al. 2006; Li 2007). The *Swift* satellite returned to observe the source at ~ 6000 s and found it in a steep decay phase with the absorbed flux ~ 100 times below the peak value. After about 1.5×10^4 s the decay of the X-ray flux flattened into a shallow afterglow with the X-ray luminosity $L_X \propto t^{-1.2}$ (Soderberg et al. 2006). The simultaneous observations by the Ultra-Violet/Optical Telescope (UVOT) onboard *Swift* satellite revealed substantial early emission peaking at about 5×10^4 s with a dereddened UV flux $\nu F_\nu \sim 10^{-10} \text{erg s}^{-1} \text{cm}^{-2}$ (Ghisellini et al. 2007a). A lower peak at a level $\nu F_\nu \approx 6 \times 10^{-12} \text{erg s}^{-1} \text{cm}^{-2}$ followed after several days. The second peak is attributed to a supernova SN2006aj (Maeda et al. 2007; Mazzali et al. 2006b; Modjaz et al. 2006; Pian et al. 2006), while the origin of the first peak is debated (Waxman et al. 2007; Ghisellini et al. 2007b,a).

The supernova was classified as Type Ic (Mazzali et al. 2006b). However, its unique spectral properties prompted to suggest a new Type Id classification (Mazzali et al. 2006b; Maeda et al. 2007). The rapid supernova peaked at 10 days (Pian et al. 2006), which is the fastest of all supernovae associated with GRBs (Ferrero et al. 2006). According to Mazzali et al. (2007), "the ejected mass predicted below 8000km s^{-1} by the model used by Mazzali et al. (2006b) is $M_{ej} = 1M_\odot$ ", while the nebular phase observations suggest a higher total ejecta mass $M_{ej} = 2M_\odot$ (Mazzali et al. 2006b, 2007; Maeda et al. 2007). The estimated ^{56}Ni mass in the ejecta is $0.2M_\odot$.

The ejecta consists mostly of oxygen and carbon (Mazzali et al. 2007), but also contains some silicon and iron (Mirabal et al. 2006). The correspondent supernova energy is $E_K \sim 2 \times 10^{51}$ erg. Both the ejecta mass and total energy are much less than the typical values $M_{ej} = 10M_\odot$ and $E_K = 3 \times 10^{52}$ erg for GRB-supernovae (see Mazzali et al. 2006b and references therein). Substantial optical polarization (Gorosabel et al. 2006) indicates the asymmetry of the ejected material.

The optical observations of the host galaxy indicate a source redshift $z = 0.0335$ (Mirabal et al. 2006; Modjaz et al. 2006), which corresponds to a distance $d = 143$ Mpc. GRB060218 is associated with a dwarf star-forming galaxy with the stellar mass $M_{st} \sim 10^{7.2}M_\odot$ (Ferrero et al. 2007), the metallicity $Z \approx 0.07Z_\odot$ (Wiersema et al. 2007), and the characteristic radius $R_{80} = 0.55$ kpc (Svensson et al. 2010). The supernova and the host galaxy were observed by *Hubble* Advanced Camera for Surveys (ACS) instrument under the program GO 10551 (PI Kulkarni) in cycle 14. The images produced by Misra et al. (2011) reveal some irregularity of the host dwarf galaxy morphology.

3. FORMER THEORETICAL MODELS

GRB060218 and the associated supernova SN2006aj sparked substantial interest among the researchers with over ten papers being dedicated to the theoretical explanations of this highly unusual source. The theoretical efforts can be divided into two big categories: a shock breakout model and a model with a jet launched by a magnetar or by a stellar mass BH. In the following subsections we briefly review the theoretical models emphasizing their successes and problems. In the following section we offer an explanation of the source within the tidal disruption scenario.

3.1. Shock Breakout Model

The shock breakout model states that the nuclear explosion following the collapse of a massive core launches a shock wave, which propagates out through the star (Colgate 1968). The radiation-dominated shock deposits its energy into the low-density gas at the stellar surface (Katz et al. 2010). The heated gas radiates bremsstrahlung photons, which are inverse-Compton scattered into the X-ray and the γ -ray bands. Gas exponentially cools down. The shock deposits up to 10^{48} erg of energy, most of which is radiated as an X-ray flash (Colgate 1974). A variety of sources were successfully modeled with the supernova shock breakout. A good example is XRF080109 and the associated Type Ib/c supernova SN2008D. The large supernova ejecta mass for that source $M_{ej} = 4 - 7M_\odot$ (Soderberg et al. 2008; Mazzali et al. 2008) and the low X-ray radiation energy $10^{45} - 10^{46}$ erg (Chevalier & Fransson 2008) leave little doubts about its shock breakout origin. Other sources include the X-ray brightening of SNLS-04D2dc supernova (Schawinski et al. 2008) with low X-ray radiated energy $\leq 10^{47}$ erg, SN1998bw with the ejecta mass $\sim 12M_\odot$ accompanied by XRF/GRB980425 (Woosley et al. 1999), SN2003lw with the ejecta mass $13M_\odot$ accompanied by a weak XRF/GRB031203 (Mazzali et al. 2006a).

In the case of XRF060218 the deposited energy $10^{49.5}$ erg is quite large (Campana et al. 2006). The peak isotropic unabsorbed flux of the soft X-ray component is $F_{BB} = (3 - 6) \times 10^{-8}$ erg $s^{-1}cm^{-2}$ with the temperature $T_{BB} = 0.11 - 0.17$ keV (Butler 2007). These parameters correspond to blackbody emission from the non-relativistic medium at a radius $R =$

$(3 - 8) \times 10^{12}$ cm, while Campana et al. (2006) inferred $R = (0.5 - 1) \times 10^{12}$ cm within their analysis at the early times and $R = (0.2 - 2) \times 10^{12}$ cm at the late times. This radius indicates a compact progenitor such as a Wolf-Rayet (WR) star. Large explosion energy might be inconsistent with the non-relativistic expansion. Mildly relativistic ejecta velocity with $v \approx 0.85c$ was proposed in later modeling by Waxman et al. (2007), who also stated that the emission should come from the photosphere above the stellar surface. The relativistic shock breakout theory was further developed by Nakar & Sari (2012), who computed the shock dynamics, the lightcurve and the spectrum. The spectrum is predicted to be quasi-thermal with the temperature around 50 keV. Nakar & Sari (2012) derived a relation between the XRF event duration, the total energy, and the observed temperature

$$t_{bo} \sim 20s \left(\frac{E_{bo}}{10^{46}\text{erg}} \right)^{1/2} \left(\frac{T_{bo}}{50\text{keV}} \right)^{-2.68}. \quad (1)$$

To satisfy this relation the early temperature of 40 keV is taken for GRB060218 consistent with the peak energy 36 keV at the very early times (Toma et al. 2007).

The shock breakout theory has its problems. Li (2007) concluded that the observed temperature and the total energy of GRB060218 lead to the unrealistically large photospheric radius R_{ph} of a WR star inconsistent with the galactic WR stellar population. This conclusion is strengthened by our larger estimate of the emission radius. The large R_{ph} in the model by Nakar & Sari (2012) corresponds to unrealistically dense stellar wind. Ghisellini et al. (2007b) showed that the optical/UV spectrum of the source is too bright for the X-rays and the optical to be the parts of the same blackbody component. In response, Waxman et al. (2007) "fiercely argued"⁸ that the shock breakout is anisotropic, and the different regions emit at the different wavelengths, which helps to reconcile the theory with the observations. The prompt thermal X-rays are emitted by a compressed shell, while the optical radiation originates in the outer shells of the expanding star at a much larger radius. In their final reply, Ghisellini et al. (2007a) found the anisotropic expansion unconvincing. The simple energetics argument is dramatically inconsistent with the shock breakout model by Waxman et al. (2007). The isotropic optical thermal emission at $t \sim 10^3$ s requires the temperature about $\sim 1 \times 10^6$ K, which corresponds to the total energy carried by a shock above 10^{51} erg, much in excess of any shock breakout model.

3.2. Jet Launched by a Magnetar or a Stellar Mass BH

A promising alternative explanation for GRB060218 is the jet launched by the central engine. A newborn central engine, a magnetar or a BH, results from a core collapse. The jet pierces through the star and escapes. That is how a typical GRB operates and our source can be just on a low-luminosity end of the spectrum. A jet with a low Lorentz factor $\Gamma \sim 5$ and a wide opening angle $\theta \sim 0.3$ can produce the emission, which peaks in the X-rays and exhibits no jet break (Toma et al. 2007). Such a jet may be powered over a long timescale by a magnetar. The central engine activity is the natural explanation of the afterglow (Soderberg et al. 2006; Fan et al. 2006). Dainotti et al. (2007) proposed a model of the interaction of an electron-positron fireshell and circumburst medium to explain

⁸ (Ghisellini et al. 2007a)

the afterglow. Their best model constrains the density profile to be $n \propto r^{-\beta}$, where $\beta = 1.0 - 1.7$ up to 10^{18} cm. The unusually bright prompt optical/UV emission can be explained as produced by the self-absorbed Comptonized synchrotron (Ghisellini et al. 2007b), while the blackbody X-rays result from the photospheric jet emission (Ghisellini et al. 2007a). Bufano et al. (2012) argued that the central magnetar would spin down too rapidly and could not power the source over thousands of seconds, while expending most of its energy in the expansion of the dense envelope. Yet, Quataert & Kasen (2012) proposed a model, in which a magnetar can power a GRB source over many days or even weeks.

The relativistic jet is a natural explanation for the source. However, if GRB060218 was a GRB, then it clearly was an unusual one. The long GRBs have large total energies $E_{\text{tot,iso}} = 10^{52-54}$ erg (Nava et al. 2008), short durations $t_{90} = 2 - 200$ s (Butler et al. 2007), and higher Lorentz factors $\Gamma = 100 - 1000$ (Pe'er et al. 2007; Liang et al. 2010; Ghisellini 2012). The high Lorentz factor of a typical long GRB is the consequence of jet confinement by the pressure from the star (Tchekhovskoy et al. 2010). Bromberg et al. (2011) argued that the GRB population does not extend to the slow low-luminosity end. Thus, another origin of events such as GRB060218 is to be sought. Below we propose that the jet is launched by a tidal disruption. This model provides the explanations for the observed event duration, the spectrum, and the lightcurve.

4. TIDAL DISRUPTION MODEL

GRB060218 can be best modeled by a low-luminosity wide mildly relativistic outflow/jet. The tidal disruption of a WD by an IMBH produces such a jet. In this section we discuss the expected temporal and spectral properties of such tidal disruptions, fit observations of GRB060218, and consider the accompanying supernova SN2006aj and the host galaxy. Extensive observations of the source pose multiple tests for any theory aimed to explain it. The tidal disruption scenario passes all these tests.

4.1. Disruption Dynamics

4.1.1. Fallback Material and Disk Formation

A star gets tidally disrupted, when it approaches sufficiently close to the BH. Let us define a tidal radius as

$$R_T = \left(\frac{2M_{\text{BH}}}{m_{\text{WD}}} \right)^{1/3} R_*. \quad (2)$$

Only the stars on orbits with the pericenter distance R_p smaller than about $R_p \lesssim R_T$ can get disrupted (see Evans & Kochanek 1989 and references therein). The WD radius is (Nauenberg 1972)

$$R_* = 8.5 \times 10^8 \left(\frac{m_{\text{WD}}}{M_{\text{Ch}}} \right)^{-1/3} \left[1 - \left(\frac{m_{\text{WD}}}{M_{\text{Ch}}} \right)^{4/3} \right]^{1/2}, \quad (3)$$

where $M_{\text{Ch}} = 1.44M_{\odot}$ is the Chandrasekhar mass. We calibrated the normalization to reproduce a typical observed WD radius (Należyty & Madej 2004) for one solar mass $m_{\text{WD}} = 1M_{\odot}$. A tidal radius depends weakly on the BH mass, while the event horizon radius $R_g = GM_{\text{BH}}/c^2$ is proportional to the BH mass. Thus, the SMBHs will swallow the WDs without disruptions and only the BHs with relatively small masses can disrupt the WDs (Luminet & Pichon 1989a; Rosswog

et al. 2009). For a star on a parabolic orbit about 50% of its material ends up being bound to the BH and about 50% remains unbound in a standard picture (Evans & Kochanek 1989). The fractions may change, when the pericenter distance is close to the BH marginally bound orbit radius (several gravitational radii). Among models with different pericenter radii, BH spins, and orientations Haas et al. (2012) found the cases, when almost all material is captured or almost all material remains unbound.

The bound material makes its way onto the BH. In a classic theory the fallback time is (Evans & Kochanek 1989)

$$t_{\text{fb}} \sim \frac{R_p^3}{\sqrt{GM_{\text{BH}}R_*^{3/2}}}. \quad (4)$$

The fallback time t_{fb} can be anywhere from several hours down to 3 min (Haas et al. 2012) for the disruptions of the WDs by the IMBHs. The peak accretion rate is achieved at

$$t_{\text{peak}} = \delta \frac{R_p^3}{\sqrt{GM_{\text{BH}}R_*^{3/2}}}, \quad (5)$$

where $\delta \approx 3.33$ for the adiabatic index $\Gamma_{\text{ad}} = 5/3$ based on simulations by Evans & Kochanek (1989) and Laguna et al. (1993) for $\beta_T = R_T/R_p = 1$, while $\delta \approx 5.5$ for $\Gamma_{\text{ad}} = 1.4$ (Lodato et al. 2009). The smaller values of the adiabatic index $\Gamma_{\text{ad}} < 5/3$ are more appropriate for the heavy WDs. The correspondent peak fallback rate is

$$\dot{M}_{\text{peak}} \approx 0.05 m_{\text{WD}} \frac{\sqrt{GM_{\text{BH}}R_*^{3/2}}}{R_p^3} \quad (6)$$

for $\Gamma_{\text{ad}} = 1.4$. The fallback rate rises from zero, when the most bound debris just reach the BH, up to the maximum value \dot{M}_{peak} at t_{peak} and then decreases according to $t^{-5/3}$ law.

The fallback matter forms an accretion disk at a circularization radius $R_{\text{circ}} = \text{several} \times R_p$. The disk exists as a radiatively inefficient accretion flow (RIAF), since the material density is very high and the photons cannot escape (Abramowicz et al. 1988). The fallback rate determines the accretion rate during the early evolution of the system. After the disk is formed, it starts to evolve on a slow viscous timescale for R_{circ} radius (Cannizzo et al. 1990). Such evolution changes the temporal slope of the accretion rate to $t^{-4/3}$ for adiabatic RIAFs (Cannizzo et al. 2011). RIAFs with the outward energy flux may unbind the material in the outer disk (Narayan et al. 2001; Metzger 2012). Then only a small percentage of the material reaches the BH. However, the inner disk regions with the radius less than $\sim 20R_g$ evolve fast and are not influenced by relatively slow energy transport (Abramowicz et al. 2002).

As was recently shown by general relativistic magneto hydrodynamic (GRMHD) simulations (Tchekhovskoy et al. 2011; McKinney et al. 2012), a poloidal magnetic field is necessary to launch a jet. The field strength for a strong jet can be estimated based on the equipartition argument. When the magnetic field energy density is comparable to the internal energy density of plasma as

$$\frac{B_{\text{BH}}^2}{8\pi} \sim 0.1 \frac{\dot{M}c}{4\pi R_g^2} \sim 0.1 n_{\text{nuc}} m_p c^2, \quad (7)$$

where n_{nuc} is the density of nucleons, m_p is the proton mass.

Then the magnetic field near the event horizon is

$$B_{BH} \sim 1.3 \times 10^{10} \left(\frac{\dot{M}}{10^3 M_\odot \text{yr}^{-1}} \right)^{1/2} \left(\frac{M_{BH}}{10^4 M_\odot} \right)^{-1} \text{ G.} \quad (8)$$

The magnetic field can reach $10^{15} - 10^{16}$ G at the base of a GRB jet (Usov 1992; Wheeler et al. 2000; Uzdensky & MacFadyen 2007; Takiwaki et al. 2009). A powerful jet is launched after a WD disruption by an IMBH for $B_{BH} \sim 10^{10}$ G. How can such a strong poloidal magnetic field be produced?

4.1.2. Generation of Magnetic Field

The magnetic field lines get advected towards the BH event horizon with the fallback material. The magnetic field of a WD is typically within $B \lesssim 10^4$ G (Angel 1978; Putney 1999) with only a few examples of a stronger field. Since the WD radius is on the order of the BH gravitational radius, then no field amplification is expected due to compression in the converging fallback flow. The BH may possess an accretion disk with a large magnetic field even before the disruption happens. When the accretion rate is close to the Eddington rate, the BH sustains the so-called Eddington magnetic field (Rees 1984; Daly 2011)

$$B_{\text{Edd}} = 6 \times 10^6 \left(\frac{M_{BH}}{10^4 M_\odot} \right)^{-1/2} \text{ G.} \quad (9)$$

The magnetic field on the order of $B \sim 6 \times 10^6$ G may already exist near the horizon of a $10^4 M_\odot$ BH. This field is higher than the advected magnetic field of a WD, but is still lower than the equipartition magnetic field for the peak accretion rate following a tidal disruption.

In situ amplification of the initial seed magnetic field into the equipartition poloidal magnetic field is needed to sustain a powerful jet. The magneto-rotational instability (MRI) (Balbus & Hawley 1991, 1998) can exponentially increase the magnetic field strength. The amplification rate of the turbulent magnetic field (to be distinguished from the instability growth rate) corresponds to a local e-folding time of 3 orbital periods (Stone et al. 1996) under the MRI action. When starting from the Eddington field strength, the magnetic field needs to grow 10^3 times to reach the equipartition. The growth by a 10^3 factor can be achieved over ~ 25 local orbits. The growth timescale is short, when the most bound material settles near the innermost stable circular orbit (ISCO) of the BH. The orbital period is (Bardeen et al. 1972; Shapiro & Teukolsky 1986)

$$t_{\text{orb}} = \frac{2\pi R_g}{c} \left[\left(\frac{R_{\text{ISCO}}}{R_g} \right)^{3/2} + a^* \right] \approx 2.5 \left(\frac{M_{BH}}{10^4 M_\odot} \right) \text{ s.} \quad (10)$$

at the ISCO for the dimensionless spin $a^* = 0.6$. In sum, a $10^4 M_\odot$ BH can amplify the random magnetic field 10^3 times over as little as 60 s.

The MRI turbulence produces the random B -field, while the regular poloidal magnetic field is needed to launch a jet. The magnetic field generation in a turbulent medium should occur via the dynamo action (Brandenburg & Subramanian 2005), which operates on a viscous timescale determined by the radial velocity v_r (Davis et al. 2010; O'Neill et al. 2011). The geometrically thick accretion flow with the effective dimensionless viscosity $\alpha = 0.1 - 0.3$ (King et al. 2007) has the radial

velocity $v_r \lesssim \alpha v_K$, where v_K is the Keplerian velocity. Then the viscous timescale at the ISCO can be as short as 15 s for a $10^4 M_\odot$ BH. Substantial generation of the regular poloidal field component is expected over $\gtrsim 10$ viscous timescales or $\gtrsim 150$ s. The failure of the previous attempts to generate the magnetic field via the dynamo action might arise from the low resolution in those simulations (Jonathan McKinney, private communication). In sum, it might be possible to generate the substantial poloidal magnetic field before the estimated accretion rate peak time.

4.1.3. Jet Launching

The jet is a relativistic outflow of the material from near the BH. The spinning BH surrounded by the accreting magnetized gas is expected to launch a jet (Blandford & Znajek 1977). We can estimate the jet power via a Blandford-Znajek formula. When the magnetic field energy density reaches the equipartition with the matter energy density close to the BH, then the Blandford-Znajek jet power is approximately (McKinney 2005; Tchekhovskoy et al. 2010)

$$L_{\text{kin}} \sim P_0 a^{*2} \dot{M} c^2. \quad (11)$$

While the simulations with the maximum magnetic flux were able to reach $P_0 \sim 1$ (Tchekhovskoy et al. 2011; McKinney et al. 2012), more traditional values observed in the simulations with the weak initial field are $P_0 = 0.01 - 0.1$ (McKinney & Blandford 2009; Penna et al. 2010). The tidal disruption of a WD by an IMBH has a typical maximum accretion rate of $\dot{M}_{\text{peak}} \sim 10^4 M_\odot \text{yr}^{-1}$ according to the formula (6). This corresponds to the jet power $L_{\text{kin}} \sim 0.01 \dot{M} c^2 = 5 \times 10^{48} \text{erg s}^{-1}$. Since there is no theoretical understanding whether or not the dynamo action would quickly generate the poloidal magnetic field, the kinetic power estimated by the equation (11) is an upper limit. The actual jet power may be substantially lower.

The kinematics of the jet is characterized by the bulk Lorentz factor Γ . The GRB jets have very high $\Gamma = 100 - 1000$ likely because of confinement by the surrounding star (Tchekhovskoy et al. 2010). In turn, the active galactic nuclei (AGN) jets are not strongly confined by the ambient gas. Their Γ -factors are about $\Gamma \sim 10$ (Jorstad et al. 2005; Pushkarev et al. 2009). Similarly, the jets from the low-mass X-ray binaries (LMXBs) have bulk Lorentz factors about $\Gamma \sim 10$ (Miller-Jones et al. 2006). The tidal disruption debris is expected to scatter in radius, cool, and provide little pressure support for the jet regardless of the direction, where the debris is scattered. Therefore, the tidal disruption jet is expected to have $\Gamma \sim 10$. A low Lorentz factor $\Gamma \sim 10$ is indeed suggested for *Swift* J1644+57 tidal disruption event (Metzger et al. 2012; Liu et al. 2012).

4.2. Prompt X-ray Emission

Prompt X-ray emission from a slow jet is expected to have a Comptonized blackbody spectrum produced near the jet photosphere. We perform time-resolved spectroscopy of *Swift* XRT observations of the source and fit the spectrum with the Comptonized blackbody model. We follow Butler (2007) and cut the full XRT window mode observations into 11 time slices with about 16,000 photons in each. We run *xrtpipeline*, select the time slices with *xselect*, take the appropriate response matrix files (RMF) from the calibration database (CALDB) (version 11) as indicated by *xrtmkarf* routine within the pipeline, and use the ancillary response files (ARF) generated by *xrtpipeline*. We run *grppha* to group

nearby bins to have at least 20 photons per bin and perform the minimization of least squares. We model the spectrum in XSPEC v12.7 (Arnaud 1996) with the broken power-law *bknpower* and with the blackbody spectrum Comptonized by thermal electrons as a part of a *compPS* model (Poutanen & Svensson 1996). We favor the exact *compPS* model over the approximate prescriptions for thermal Comptonization of the blackbody spectrum such as *compBB* (Nishimura et al. 1986) and *compTT* (Titarchuk 1994). We compute the model flux in (0.05–10) keV band to capture very soft thermal X-rays. Unlike Butler (2007) and Campana et al. (2006), we consider the host galaxy to be a low-metallicity absorber with $Z = 0.07Z_{\odot}$ (Wiersema et al. 2007), while following Campana et al. (2006) we fix the Galactic column at $N_H = 9.4 \times 10^{20} \text{cm}^{-2}$ with the solar metallicity. We ignore any potential changes of the hydrogen absorption column with time. We search for a joint best fit to all time slices with a single host galaxy N_H .

The results of spectral fitting are presented in Table 1 for the broken power-law fits. A substantial degeneracy exists between the soft power-law slope in a *bknpower* model and the host N_H . We fix the soft photon index to be $\Gamma_1 = -1$, which is representative of the Rayleigh-Jeans tail $F_{\nu} \propto \nu^2$ of the blackbody. The host galaxy hydrogen column $N_H = 0.750 \times 10^{22} \text{cm}^{-2}$ provides the best joint fit to all time slices with $\chi^2 = 5463.0$ for dof = 4836 degrees of freedom. The break energy is 0.7–0.8 keV and depends weakly on the total flux. Since the emission becomes substantially softer with time, the absorbed flux constitutes a much smaller fraction of the unabsorbed flux at the late times. This masks the true temporal evolution of the source flux.

The results of spectral fitting are presented in Table 2 for the blackbody spectrum Comptonized by thermal electrons as computed within a *compPS* XSPEC model. The optical depth of Compton scattering is fixed at $\tau_{\sigma} = 1$ for scattering at the photosphere. The host galaxy hydrogen column $N_H = 1.090 \times 10^{22} \text{cm}^{-2}$ provides the best joint fit to all time slices with $\chi^2 = 5358.9$ for dof = 4836 degrees of freedom. The fit is slightly better for this model compared to the *bknpower* model. Our fit has a higher blackbody flux fraction at the peak (40%) compared to $\sim 15\%$ fraction in a fit by Campana et al. (2006), who employed the solar metallicity of the absorber and used a sum of the power-law and the blackbody. Table 2 also summarizes application of the fireball model (Pe'er et al. 2007) to the blackbody emission component. We can estimate the bulk Lorentz factor of the flow Γ and the jet base radius knowing the temperature T_{BB} , the observed isotropic flux F_{BB} , the distance to the source, and assuming a fiducial ratio $Y = 10$ of the total fireball energy to the energy emitted in X-rays/ γ -rays. The obtained jet parameters are surprisingly consistent between different time slices, which might indicate that the fireball model captures the physics of the event. The bulk Lorentz factor of $\Gamma \approx 2.7Y_{10}^{1/4}$ corresponds to a mildly relativistic jet, while the jet base radius is $R_0 = (1.0-1.7) \times 10^{10} Y_{10}^{-3/2} \text{cm}$. The model is self-consistent for a wide range of Y as the saturation radius $R_s = \Gamma R_0$ lies well within the photosphere with the radius $R_{\text{ph}} \sim 5 \times 10^{12} \text{cm}$. The jet launching region should be located at several BH gravitational radii. Assuming that the jet base radius is at $R_0 = 5R_g$, we find the BH mass $M_{\text{BH}} = (1.3-2.3) \times 10^4 Y_{10}^{-3/2} M_{\odot}$, which places the BH into the intermediate mass category. Pe'er et al. (2007) argue that the ratio of energies is $Y_{10} \lesssim 1$.

The reduced chi-squared $\chi^2/\text{dof} = 1.138$ for the joint *bknpower* fit and $\chi^2/\text{dof} = 1.115$ for the joint *compPS* fit are quite low despite good photon statistics. Even better fits down to $\chi^2/\text{dof} = 1.0$ can be achieved, when the host galaxy hydrogen column density is allowed to vary between the time slices as in Butler (2007). However, we do not see a physical reason for N_H to vary on the timescale of 10^3s , which corresponds to $3 \times 10^{13} \text{cm}$ distance, while the characteristic star-formation size of the host galaxy is $R_{80} = 0.55 \text{kpc} = 10^{21} \text{cm}$ (Svensson et al. 2010). To test for potential discrepancies between the fits and the observations, we depict on Figure 1 the parts of fits and the normalized residuals for the 4-th time slice characterized by the highest soft X-ray flux. The left panel shows the results for the *bknpower* model, while the right panel shows the results for the blackbody spectrum Comptonized by thermal electrons. The residuals do not systematically deviate from zero except slightly at the very low energies or for certain absorption features. At the low energies *Swift* XRT response might be uncertain or the line physics might alter the spectrum. The most prominent absorption feature is around 7.3 keV.

4.3. Prompt Optical/UV Emission

The prompt optical/UV emission observed by the UVOT instrument onboard *Swift* satellite was a matter of major disagreements between the former theoretical models. The same blackbody component cannot produce both the optical and the X-ray emission (Ghisellini et al. 2007b), which is consistent with the physical picture within the tidal disruption scenario. For an estimate we take the magnetic field in the photosphere to be in equipartition with the jet kinetic energy. Then the Poynting energy flux is about $Y/2 \sim 5$ times larger than the total X-ray radiation flux. Hence the magnetic field in the photosphere has the strength

$$B_{\text{ph}} \sim 10^5 \text{G}. \quad (12)$$

The *compPS* fit to the prompt X-ray spectrum shows that the non-thermal particles with the energies $\sim 200 \text{keV}$ exist in the photosphere. The fact that the emission continues to the hard X-rays in BAT band reveals the existence of much more energetic particles with $\gamma \sim 10-100$, which readily emit synchrotron in the optical band. The synchrotron emission by an electron with a random Lorentz factor $\gamma \sim 30$ peaks at 300 nm for 10^5G magnetic field. Since the effective synchrotron cross-section is much higher than the Thompson scattering cross-section, the optical photosphere is much above the X-ray photosphere. In fact, very large jet particle density makes the optical photosphere coincide with the front of the jet propagating with a bulk Lorentz factor $\Gamma \sim 3$ into the surrounding medium. The jet front can propagate up to $R_{\text{front}} = t_{\text{obs}} \Gamma^2 c$, where t_{obs} is the time since BAT trigger. Continuous energy dissipation and particle acceleration at the front ensures there are enough energetic electrons to emit optical/UV synchrotron. Note that the interactions with the interstellar medium (ISM) at the early times, which lead to a forward shock, may contribute less to the emission. Low magnetic field at the jet front

$$B_{\text{front}} \sim 10^3 \text{G} \quad (13)$$

at $t_{\text{obs}} = 3 \times 10^3 \text{s}$ is compensated by higher Lorentz factors $\gamma \sim 300$ of the electrons producing the optical/UV emission.

The idea of the jet front emission explains the steep $F_{\nu} \propto \nu^2$ spectrum and the achromatic flux growth at different UVOT

TABLE 1
TIME-RESOLVED SPECTROSCOPY OF GRB060218 SOFT X-RAY SPECTRUM. FITTING WITH A BROKEN POWER-LAW.

Number	Time period t [s]	Break energy E_{break} [keV]	High energy slope Γ_2	Absorbed flux $F_{\text{abs}} [10^{-9} \text{erg s}^{-1} \text{cm}^{-2}]$	Unabsorbed source flux $F_{\text{unabs}} [10^{-9} \text{erg s}^{-1} \text{cm}^{-2}]$
1	164-478	0.795	1.226	4.753	5.234
2	478-691	0.740	1.171	7.411	8.148
3	691-875	0.820	1.234	8.242	9.064
4	875-1049	0.746	1.290	8.572	9.575
5	1049-1226	0.813	1.419	7.595	8.575
6	1226-1414	0.800	1.519	6.780	7.795
7	1414-1620	0.761	1.650	5.583	6.631
8	1620-1854	0.749	1.816	4.421	5.467
9	1854-2119	0.678	1.962	3.621	4.783
10	2119-2404	0.712	2.190	2.940	4.110
11	2404-2756	0.656	2.332	2.265	3.457

NOTE. — The parameters of XSPEC *bknpower* model are shown. The low energy slope of the count spectrum is fixed at $\Gamma_1 = -1$, which corresponds to the Rayleigh-Jeans tail of the blackbody emission with $F_\nu \propto \nu^2$. The host galaxy hydrogen column density $N_H = 0.750 \times 10^{22} \text{cm}^{-2}$ with the metallicity 0.07 of solar (Wiersema et al. 2007) provides the best joint fit to all time slices with $\chi^2 = 5463.0$ for dof = 4836. The galactic column is fixed at $N_H = 9.4 \times 10^{20} \text{cm}^{-2}$ following Campana et al. (2006).

TABLE 2
TIME-RESOLVED SPECTROSCOPY OF GRB060218 SOFT X-RAY SPECTRUM. FITTING WITH THE BLACKBODY SPECTRUM COMPTONIZED BY THERMAL ELECTRONS.

Number	Time period t [s]	Photon temperature T_0 [keV]	Electron temperature T_e [keV]	Absorbed flux $F_{\text{abs}} [10^{-9} \text{erg s}^{-1} \text{cm}^{-2}]$	Unabsorbed source flux $F_{\text{unabs}} [10^{-9} \text{erg s}^{-1} \text{cm}^{-2}]$	Blackbody source flux $F_{\text{BB}} [10^{-9} \text{erg s}^{-1} \text{cm}^{-2}]$	Lorentz factor $\times Y_{10}^{1/4}, \Gamma$	jet base radius $\times Y_{10}^{-3/2}, R_0 [10^{10} \text{cm}]$
1	164-478	0.1039	262.3	4.586	6.61	2.630	2.667	0.906
2	478-691	0.1052	297.6	7.159	10.17	4.033	2.832	1.091
3	691-875	0.0994	258.8	7.965	11.49	4.465	2.803	1.245
4	875-1049	0.1058	230.2	8.353	12.29	5.020	2.898	1.257
5	1049-1226	0.0967	176.6	7.424	11.52	4.763	2.744	1.493
6	1226-1414	0.1008	145.8	6.636	10.84	4.848	2.753	1.560
7	1414-1620	0.1119	119.5	5.517	9.56	4.784	2.816	1.487
8	1620-1854	0.1116	93.1	4.369	8.43	4.686	2.732	1.733
9	1854-2119	0.1324	80.9	3.614	7.16	4.412	2.879	1.393
10	2119-2404	0.1489	65.9	2.970	6.25	4.284	2.961	1.275
11	2404-2756	0.1421	55.3	2.268	5.41	3.963	2.818	1.485

NOTE. — The parameters of the blackbody spectrum Comptonized by thermal electrons within a *compPS* model are shown. The optical depth is fixed at $\tau_\sigma = 1$, which corresponds to Compton scattering at the photosphere. We find the host galaxy hydrogen column density $N_H = 1.090 \times 10^{22} \text{cm}^{-2}$ in a self-consistent joint fit over all time slices with $\chi^2 = 5358.9$ for dof = 4836 degrees of freedom. The Galactic column is fixed at $N_H = 9.4 \times 10^{20} \text{cm}^{-2}$ with solar metallicity following Campana et al. (2006). The quantity Y_{10} is the ratio of the total fireball energy to the energy emitted in the X-rays in the units of 10. The last two columns are computed for $Y_{10} = 1$ or $Y = 10$ within the thermal fireball emission model described in Pe'er et al. (2007). The jet base radius $R_0 \sim 10^{10} \text{cm}$ corresponds to the BH mass $M_{\text{BH}} \sim 10^6 M_\odot$ assuming the jet is launched within several R_g from the center.

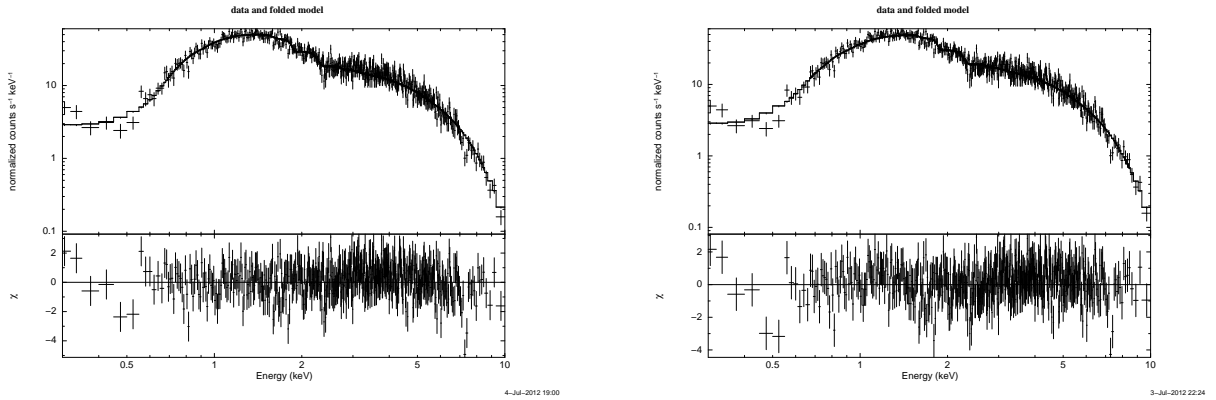


FIG. 1.— The parts of XSPEC spectral fits and the normalized residuals for the 4-th time slice. This time slice corresponds to the highest soft X-ray flux. The left panel shows the results for the *bknpower* model, the right panel shows the results for the blackbody spectrum Comptonized by thermal electrons within the *compPS* model. The model parameters can be found in Tables 1 and 2.

frequencies till $\sim 3 \times 10^4$ s. The jet front is transparent to the hard UV and X-ray photons. Since the particles are subject to rapid cooling, the spectrum of re-accelerated particles may substantially deviate from a power-law and resemble a very hot Maxwellian. Then the resultant self-absorbed spectrum could be the Rayleigh-Jeans part of the blackbody $F_\nu \propto \nu^2$ as opposed to the absorbed synchrotron $F_\nu \propto \nu^{5/2}$ spectrum (Rybicki & Lightman 1979). The observed temperature of the optical/UV radiation is about 10^6 K (Ghisellini et al. 2007a), which violates the energetics of the source in a non-relativistic emission model with the total emitted energy $E_{BB} \gtrsim 10^{51}$ erg. In turn, the relativistic jet front emission does not violate the source energetics. Relativistic Doppler boosting with $\Gamma \sim 3$ reduces the emission temperature down to $\lesssim 3 \times 10^5$ K, which readily leads to $E_{BB} \lesssim 10^{50}$ erg. The inferred large optical emission radius 10^{14-15} cm at 3×10^3 s is consistent with R_{front} for the mildly relativistic expanding shell. A more detailed jet model for optical/UV emission by Ghisellini et al. (2007b) is similarly consistent with the source energetics.

4.4. Temporal Analysis

The spectral modeling allows for proper determination of fluxes at different times to study the temporal behavior of the source. Figure 2 shows the absorbed and the unabsorbed source fluxes as functions of time. The soft X-ray flux peaks at around 800 s following *Swift* BAT trigger. We fit the lightcurve with the scaled accretion rate dependence on time $\dot{M}(t)$ for a tidal disruption of a MS star by a SMBH. Such a dependence for $\beta_T = 1$ was taken from Laguna et al. (1993). The best fit corresponds to a time delay of $\Delta t = 1810$ s between the tidal disruption and *Swift* BAT trigger. The peak at $t_{\text{peak}} \approx 2600$ s corresponds, for example, to the tidal disruption of a $m_{WD} = 0.75M_\odot$ WD by a $1 \times 10^4 M_\odot$ IMBH (or a $m_{WD} = 0.86M_\odot$ WD by a $2 \times 10^4 M_\odot$ IMBH) at $\beta_T = 1$. The tidal radius in this case is $16R_g$ (or $8R_g$). A deeper encounter with $\beta_T > 1$ may be needed to initiate a nuclear burning for an accompanying supernova, but the dependence of critical β_T , for which the effective nuclear burning starts, on the WD and the BH masses is not known at present. The encounter in question may indeed have very small pericenter distance with $\beta_T \gg 1$ and not violate the peak time constraint, as De Colle et al. (2012); Guillochon & Ramirez-Ruiz (2012) find very weak dependence of peak times on β_T .

The accretion rate curve $\dot{M}(t)$ provides a good fit to the unabsorbed source flux. However, such fitting procedure is prone to multiple caveats. First, the accretion disk onto the BH takes a finite amount of time $t_{\text{magn}} \sim 100t_{\text{orb}} \sim 250$ s to generate the strong poloidal magnetic field. Despite the magnetic field generation time is much less than the characteristic dynamical time $t_{\text{magn}} \ll t_{\text{dyn}} \sim 1000$ s, the precise fit is not expected. Secondly, the transition from the accretion power $\dot{M}c^2$ to the emission power in a certain band is not trivial. The jets are expected to be radiatively efficient (Pe'er et al. 2007), and most of the emission, especially at the late times, falls into the 0.05–10 keV band. Only if the jet kinetic power is a constant fraction of the accretion power, then the soft X-ray luminosity is expected to approximately follow the accretion power. Finally, the GR effects are not included in these estimates. Haas et al. (2012) emphasized that the ultra-close encounters with small $R_p \sim \text{several} \times R_g$ may lead to non-trivial temporal behavior of \dot{M} .

The estimated large BH mass and the small Lorentz factor

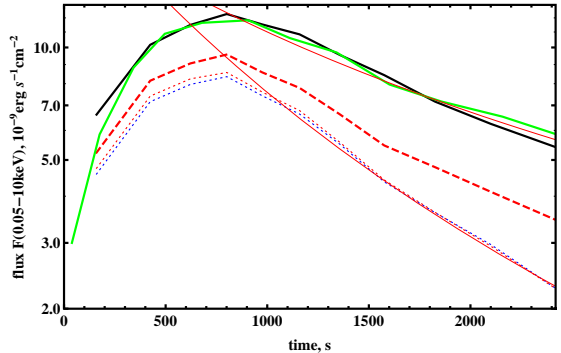


FIG. 2.— Soft X-ray lightcurve of GRB060218 fitted with the accretion rate temporal evolution $\dot{M}(t)$. Shown are the unabsorbed source lightcurve for the *compPS* model (dark solid) and for the *bknpower* model (red dashed), the observed absorbed lightcurve for the *compPS* model (lower dotted) and for the *bknpower* model (upper dotted), and the scaled fallback accretion rate based on Laguna et al. (1993) (light solid). The scaled accretion rate is offset by $\Delta t = 1810$ s, which the most bound debris takes to reach the BH. Two thin lines $(t + \Delta t)^{-3}$ and $(t + \Delta t)^{-5/3}$ are shown to emphasize the asymptotic behavior of the absorbed flux and the unabsorbed flux, respectively. While the observed flux decays as $F \propto t^{-3}$, the unabsorbed flux exhibits $F \propto t^{-5/3}$ late-time behavior in agreement with the tidal disruption scenario.

may lead to a smooth lightcurve. The lightcurves of typical GRBs are highly variable, which is attributed to fast variations near small BHs with masses $M_{BH} \sim 10M_\odot$, when the BHs launch shocks at different speeds. The variation timescale is correspondingly longer for the BHs with much larger masses $M_{BH} \sim 10^4 M_\odot$. Those variations amplify, when the shocks moving at different high Lorentz factors collide. The energy release and the radiated energy are relatively small, when the shocks with small Γ 's collide. Lastly, the collisions of the slow-moving shocks are likely to happen inside the photosphere, so that any variations are smoothed. Thus, a slow moving jet of GRB060218 is expected to vary on a very long timescale and the variations are expected to be small. The object should exhibit a smooth lightcurve. On Figure 3 we

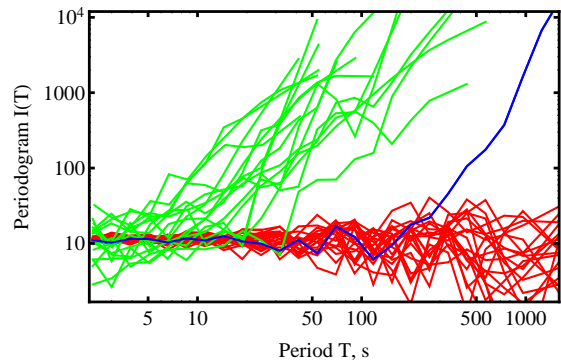


FIG. 3.— Log-smoothed periodograms $I(T)$ (Papadakis & Lawrence 1993) of the lightcurves binned to 1 s. The flat red curves depict simulated white noise, GRB060218 is shown by a dark/blue curve, while 16 other GRBs are shown by light/green curves. The white noise has the same mean count rate and total counts as GRB060218. The other GRBs are selected to have the peak count rate between 142cts s^{-1} and 166cts s^{-1} , while the peak count rate is 148cts s^{-1} for GRB060218. The break time t_b , at which the periodogram starts deviating substantially from the noise, corresponds to a characteristic system timescale. The break time for our event is $t_b \sim 300$ s, which is consistent with the emission rise timescale in the tidal disruption. Other GRBs have substantially lower break times $t_b = 5 - 40$ s, and the distribution of t_b does not continue to 300 s. This indicates a different origin of GRB060218. Log-smoothing is performed to 0.11 dex.

present the log-smoothed periodograms $I(T)$ (Papadakis &

Lawrence 1993) of GRB060218 and other lightcurves. The dark (blue) rising curve is the periodogram of GRB060218 prompt X-rays. The dark (red) constant curves are the periodograms of 20 white noise implementations with the same mean count rates and total counts as GRB060218. The light (green) curves represent the selection from the online *Swift* catalogue of 16 GRBs, whose peak photon count rates are the closest to the peak count rate 148cts s^{-1} of GRB060218. The white noise implementations are flat as expected, while our candidate tidal disruption source rises above the noise at a very long break time of about $t_b = 300$ s. Other 16 GRBs show much smaller break times of 5–40 s. For some of the sources the break time coincides with the characteristic flux decay time, while most sources exhibit variability on a shorter timescale. As predicted, the lightcurve of GRB060218 is smooth and shows no variability faster than the emission rise timescale. This strengthens the case for the source being a slow dense jet from an IMBH.

4.5. Steep Decay Phase and Afterglow

The *Swift* satellite was unable to observe the source from 2780 s till 5900 s counting from the BAT trigger. The observed X-ray flux at 5900 s is very low, which is seemingly inconsistent with $t^{-5/3}$ behavior of the accretion rate. However, the blackbody source flux may still be large, while the observed flux may be a factor of 40 lower due to the action of absorption and incomplete overlap with the XRT band. The XRT instrument observes mostly Comptonized photons with its effective energy range 0.3–10 keV. The energy of the Comptonizing electrons goes down with time during the prompt phase, which translates into the ratio of unabsorbed to absorbed fluxes rising from 1.4 to 2.4 by 11-th time slice (see Table 2). The ratio of fluxes keeps dropping till 10^4 s in the steep decline phase. After that a much shallower decline $F \propto t^{-1.2}$ follows (Campana et al. 2006) with the spectrum consistent with a power-law $F_\nu \propto \nu^{-2.2}$ (Soderberg et al. 2006). While the temporal decay of this late phase is consistent with the afterglow, the spectrum is substantially softer than the typical afterglow spectrum emitted by the external shock (Toma et al. 2007). The softer spectrum and the long afterglow may be explained by the late activity of the source (Soderberg et al. 2006). This is an especially viable idea, since the jet kinetic luminosity is expected to stay high for a long period of time following a tidal disruption.

We model the steep decay spectrum within the time interval $t = (5950, 7070)$ s with a combination of the blackbody component Comptonized by thermal electrons and the power-law emission from the external shock. We fix the absorption at the level determined with *compPS* fitting of the prompt X-ray emission. The fit allows for a large range of power-law slopes without a significant change in χ^2 . We fix the power-law slope at $\Gamma_{pl} = 2$ ($F_\nu \propto \nu^{-1}$), which corresponds to the afterglow of a typical GRB (Soderberg et al. 2006). The absorbed combination of the *compPS* and the *powerlaw* models leads to $\chi^2 = 43.1$ for dof = 36. As suggested by Butler (2007), the spectrum contains lines, which become prominent at the late times. Lines is the natural outcome of the atomic processes in a warm absorbing/emitting jet baryonic material. The jet baryonic material consists of WD debris rich in oxygen and carbon and practically devoid of hydrogen. The absorption by cooled down jet exhaust may become strong at the late times. We model such absorption by a blueshifted *zlyphabs* model, where we set the abundance of a single chem-

ical element to 1000 and the abundances of the other elements to zero. The model with a pure oxygen cold absorber converges to a blueshift $z = -0.422$, which corresponds to the bulk Lorentz factor $\Gamma = 1.73$ consistently with the fits to the prompt spectrum. The fit with $\chi^2 = 37.3$ for dof = 36 and the normalized residuals are depicted on Figure 4. The fit can be marginally improved with the addition of iron elements. We do not present such modeling here for simplicity. The addition of up to 30% of carbon does not change the fit. The best-fitting oxygen column density $N_O = 1 \times 10^{18}\text{cm}^{-2}$

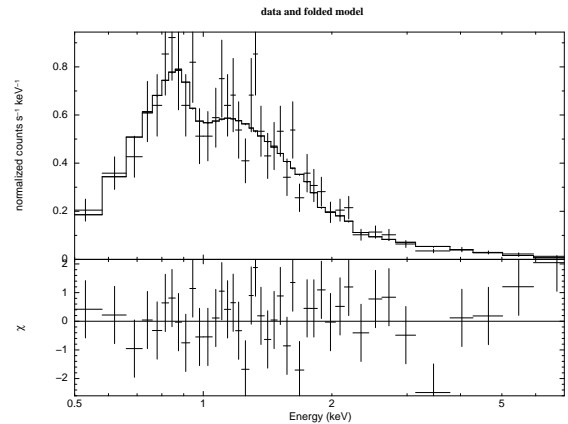


FIG. 4.— Best fit to the spectrum at the late times $t = (5950, 7070)$ s with $\chi^2 = 37.3$ for dof = 36. The model consists of the blackbody emission Comptonized by thermal electrons absorbed by blueshifted oxygen and a powerlaw with a fixed slope $\Gamma_{pl} = 2$ ($F_\nu \propto \nu^{-1}$). The sum of these two components is absorbed by the galactic $N_H = 9.4 \times 10^{20}\text{cm}^{-2}$ and $N_H = 1.090 \times 10^{22}\text{cm}^{-2}$ of the host galaxy with their respectively metallicities all fixed from the early time modeling. The blueshift $z = -0.422$ corresponds to the bulk Lorentz factor $\Gamma = 1.73$. The best-fitting cold oxygen column density is $N_O = 1 \times 10^{18}\text{cm}^{-2}$, which can readily be provided by a cooling jet exhaust. The blueshifted oxygen manifests as an absorption feature around 1 keV. The unabsorbed source flux in $F_{\text{unabs}}(0.05 - 10\text{keV}) = 1.8 \times 10^{-9}\text{erg s}^{-1}\text{cm}^{-2}$, which coincides with the extrapolation based on $t^{-5/3}$ law, while the heavily absorbed observed flux is $F_{\text{abs}} = 4.0 \times 10^{-11}\text{erg s}^{-1}\text{cm}^{-2}$.

can be readily provided by the cooling jet material. Very soft blackbody has a temperature $T = 0.052$ keV. The best-fitting flux is $F_{\text{unabs}}(0.05 - 10\text{keV}) = 1.8 \times 10^{-9}\text{erg s}^{-1}\text{cm}^{-2}$, which directly traces $t^{-5/3}$ decay law from the time of peak flux, while the observed absorbed flux $F_{\text{abs}} = 4.0 \times 10^{-11}\text{erg s}^{-1}\text{cm}^{-2}$ is 40 times lower. The absorbed power-law contribution to the total source flux is $F_{\text{abs,pl}} = 2.5 \times 10^{-11}\text{erg s}^{-1}\text{cm}^{-2}$, which is consistent with the emission at later times being dominated by the power-law. The application of the photospheric emission model (Pe’er et al. 2007) to determine the Lorentz factor and the jet base radius gives $\Gamma = 1.45Y_{10}^{1/4}$ and $R_0 = 1.1 \times 10^{11}Y^{-3/2}$ cm, respectively. A larger ratio of total to radiated energy $Y = 4 \times Y_{\text{prompt}}$ makes the numbers consistent with the results from the prompt phase. Since the blackbody is too soft to be directly observed, there is a substantial degeneracy between its temperature and normalization, which leads to the unabsorbed flux being uncertain by a factor of 3. Taking this uncertainty into the account, the unabsorbed flux F_{unabs} is consistent with the $t^{-5/3}$ law even for larger Y . We might not be able to compute a more self-consistent physical model of the steep decay phase unless we incorporate the atomic physics of warm emitters/absorbers, which lies beyond the scope of the present paper.

According to Cannizzo et al. (2011) switching from $t^{-5/3}$ fallback dominated behavior to $t^{-4/3}$ disk spreading behavior of the accretion rate occurs at

$$t_x = 1.25 \times 10^{10} \text{s} \left(\frac{M_\star}{M_\odot} \right)^{-3/2} \left(\frac{R_\star}{R_\odot} \right)^{3/2} \left(\frac{M_{\text{BH}}}{10^7 M_\odot} \right) \left(\frac{\alpha}{0.1} \right) \left(\frac{\eta}{2\beta_T} \right)^{9/2}, \quad (14)$$

where $\eta = R_{\text{circ}}/R_P$ is the effective extent of the disk. Taking a WD with a mass $0.86(0.75)M_\odot$ and a radius $R_\star = 7.1(8.1) \times 10^8$ cm disrupted by a $2.0(1.0) \times 10^4 M_\odot$ BH we estimate the transition time to be $t_x = 2.6(1.6) \times 10^4 \beta_T^{-9/2}$ s for $\eta = 2$. Thus, the transition to the shallow disk spreading behavior occurs around the transition to the afterglow phase. The temporal index of the power-law flux decay -1.2 ± 0.1 during the afterglow (Soderberg et al. 2006) is consistent with -1.33 index of the mass accretion rate at the late times.

4.6. Associated Supernova and Host Galaxy

Another challenge to the WD/IMBH tidal disruption model is the associated supernova. As discussed in § 2, the supernova ejecta mass $M_{\text{ej}} \sim (1-2)M_\odot$ is consistent with the high end of the WD mass distribution. The unbound mass fraction is influenced by the interplay between the energy spread during the tidal disruption and the energy release in the supernova explosion. The outflow velocity of tidal disruption debris is $20 \times 10^3 \text{km s}^{-1}$ for the WD and the BH masses estimated above, which is comparable to the typical outflow velocity in a Type Ia supernova $(10-20) \times 10^3 \text{km s}^{-1}$ (Wang et al. 2009). Thus, the supernova SN2006aj with the energy release comparable to that of a Type Ia supernova (Khokhlov et al. 1993) can unbind most of the WD material, which then contributes to the supernova ejecta mass. Only a small fraction of debris accretes back onto the BH. The nebular phase of SN2006aj exhibits strong oxygen emission (Mazzali et al. 2007), which is consistent with the ignition of a carbon/oxygen WD.

A supernova explosion accompanying a tidal disruption is expected to be heavily asymmetric. In particular, there could be a velocity shift of supernova lines with respect to the velocities of the host galaxy. Yet, the velocities and the redshift $z \approx 0.033$ were only estimated for the host galaxy (Mirabal et al. 2006; Modjaz et al. 2006) and no measurements of supernova velocities exist.

Since the IMBHs are expected to reside in the nuclei of dwarf galaxies, the other major test is the coincidence of the supernova position with the center of the host galaxy. A star-formation radius for GRB060218 host is $R_{80} = 0.55$ kpc (Svensson et al. 2010), which corresponds to a half-light radius $R_{50} \approx 0.45$ kpc. The supernova and the host galaxy were observed with *Hubble* ACS instrument at different epochs. Figure 5 shows the position of the supernova (green contours) relative to the intensity plot of the host galaxy at a late time. The host galaxy has a somewhat irregular morphology as can be seen on images by Misra et al. (2011), which complicates finding its center. The uncertainty of center determination is about $dR \lesssim 0.1 \text{arcsec} \approx 70$ pc. The supernova position appears to be consistent with the center to within 0.1arcsec . The ratio of the uncertainty dR to the half-light radius is $dR/R_{50} < 0.15$, which is much less than a normalized offset $dR/R_{50} = 0.98$ found for a representative set of long GRBs (Bloom et al. 2002). It is very rare for a supernova to happen that close to a galactic center at random.

The stellar mass of the host galaxy is estimated to be $M_{\text{st}} \sim 10^{7.2} M_\odot$ (Ferrero et al. 2007). Observations suggest

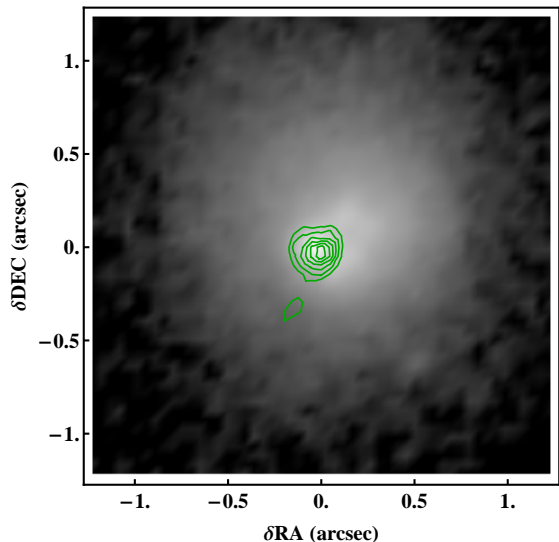


FIG. 5.— Position of SN2006aj supernova (green contours) within the host galaxy (grayscale intensity plot) from *Hubble* ACS observations in red F814W band. The supernova position is consistent with the center of the host. The supernova image is computed as the subtraction of the late time (26 Nov 2006) image from the early time (18 Jul 2006) image. The late time image without supernova subtraction is taken to represent the host galaxy, which has irregular morphology. The subtraction is not performed to avoid the oversubtraction of the galactic nucleus were the supernova to happen at the very center.

that the bulge mass to the total stellar mass ratio (B/T) of galaxies is within 15–100% range (Schramm & Silverman 2012) and $B/T = 15\%$ for the Milky Way (McMillan 2011). Then, extrapolating $M_{\text{bulge}} - M_{\text{BH}}$ relation (Magorrian et al. 1998; Marconi & Hunt 2003)

$$M_{\text{BH}}/M_{\text{bulge}} \sim 0.002, \quad (15)$$

to low masses, we get the central BH mass to be $M_{\text{BH}} \sim (0.5-3) \times 10^4 M_\odot$. This estimate, though much less certain, is consistent with two previous estimates based on the jet thermal emission and the lightcurve peak time.

4.7. Accretion Flow and Jet Energetics

Let us compute the energetics of the jet for the estimated BH mass. Assuming the supernova does not unbind the gas, the accretion rate onto the BH peaks at

$$\dot{M} \approx 2 \times 10^3 M_\odot \text{yr}^{-1}, \quad (16)$$

which corresponds to the accretion power $\dot{M}c^2 \approx 1 \times 10^{50} \text{erg s}^{-1}$. Then the Blandford-Znajek jet power is $P_{\text{BZ}} \sim 0.01 \dot{M}c^2 \sim 1 \times 10^{48} \text{erg s}^{-1}$ (McKinney 2005) for the low efficiency associated with simulations initiated with the weak magnetic field. The isotropic equivalent radiation power for the opening angle $\theta = 0.5$ and the ratio $Y = 10$ is $P_{\text{rad,iso}} \sim 2 \times 10^{48} \text{erg s}^{-1}$. This is substantially higher than the observed peak isotropic luminosity $P_{\text{obs,iso}} = 3 \times 10^{46} \text{erg s}^{-1}$. The discrepancy can be partially explained by a small fraction of a WD material falling back onto the BH, since the supernova may unbind most of the gas. Other potential explanations include GRB060218 happening substantially off-axis, the WD being mostly unbound in an encounter with a spinning BH (Haas et al. 2012), or the source failing to launch a powerful jet. The latter is especially viable, since the source has a very limited amount of time to generate the ordered magnetic field. A poloidal magnetic field 10 times weaker than the equipartition value given by the equation (8) leads to consistent jet

energetics. The disk may launch a mildly relativistic outflow by itself via the Blandford-Payne mechanism (Blandford & Payne 1982). An off-axis GRB is inconsistent with the absence of a jet break (Soderberg et al. 2006).

The X-ray emission is dominated by the super-Eddington jet at the early times, but the accretion disk may be brighter at the late times. Our disruption models have the peak accretion rate of about $\dot{M}_{\text{peak}} \approx 2 \times 10^3 M_{\odot} \text{yr}^{-1}$ at the time 2600 s after the disruption (neglecting the effect of blowing the material away by a supernova). Then assuming $t^{-5/3}$ law we find that the accretion rate is down to the Eddington value for a standard $\varepsilon = 0.1$ efficiency at $t_{\text{Edd}} \sim 1$ yr. As described in § 4.5, the accretion rate is expected to follow a shallow $t^{-4/3}$ slope after about $t_x \sim 2 \times 10^4$ s due to the viscous evolution in the disk. Yet, as we will describe below in § 5.1, when the outer disk boundary diffuses out to be much larger than $\sim 20R_g$, then the energy transfer between the inner and the outer disk kicks in. The energy from the inner flow unbinds the outer material and the accretion rate onto the BH dramatically decreases. Cannizzo & Gehrels (2009) derived the temporal dependence of the accretion rate in the RIAF with the energy transfer/convection to be $t^{-4(1+s)/3}$, where the parameter s controls the convection efficiency. The value of s varies from $s = 0$ for no energy transfer up to $s = 1$ for the maximum energy transfer power. Then the accretion rate may decrease as steeply as $t^{-8/3}$ in the RIAF phase. We will discuss the RIAFs with convection in more detail in § 5.1. The onset of RIAF with the energy transfer may explain the steeper XRT flux dependence $F \propto t^{-1.55}$ at times $t > 2 \times 10^5$ s, accompanied by a nearly constant hardness ratio. Therefore, the accretion rate may fall below the Eddington rate at times $t \ll 1$ yr.

Yet, the Eddington luminosity for a $2 \times 10^4 M_{\odot}$ BH corresponds to a flux $F_X \sim 10^{-12} \text{erg s}^{-1} \text{cm}^{-2}$, which is about a factor of 100 higher than the late-time flux observed by the XRT. This is not a contradiction, however, as at times much before 1 yr the accretion disk might not be visible at all due to obscuration by the outflowing debris. As discussed in our previous paper (Haas et al. 2012), the outflowing debris consisting of carbon and oxygen have high absorption cross-section $\sigma \approx 2 \times 10^{-20} \text{cm}^2$ across 0.5–10 keV band. The observed matter velocity in SN2006aj is about $1.5 \times 10^4 \text{km s}^{-1}$ (Mazzali et al. 2006b). A tidal disruption may scatter matter in particular directions leaving huge voids, though which the inner disk could be observed. A tidal disruption of an object with the size comparable to the tidal radius rather tends to scatter the debris over the large solid angle (Haas et al. 2012). In addition, the supernova, whose energy release is comparable to the energy release in a tidal disruption, should scatter the debris more uniformly in all directions. Following Haas et al. (2012) we find that the absorption optical depth unity $\tau = 1$ is achieved around $t \sim 3$ yrs, at which time the source is expected to rebrighten. As the accretion disk should have a substantially sub-Eddington accretion rate by 3 yrs, the emitted flux levels $F_X \lesssim 10^{-13} \text{erg s}^{-1} \text{cm}^{-2}$. The accretion flow with the sub-Eddington accretion rate settles into a thin disk (Shakura & Sunyaev 1973), which has the inner temperature about

$$T_{\text{in}} \approx \left(\frac{GM_{\text{BH}}\dot{M}}{\sigma_F R_{\text{ISCO}}^3} \right)^{1/4}. \quad (17)$$

The temperature is $T_{\text{in}} \approx 0.15$ keV for the accretion rate equal to 10% of the Eddington value and the BH mass

$2 \times 10^4 M_{\odot}$. Very soft blackbody spectrum absorbed by the galactic column and the host galaxy column leads to the ratio $F_{\text{unabs}}/F_{\text{abs}} \sim 10$ of the model flux to the observed absorbed flux and the observed flux $F_{\text{abs}} \sim 10^{-14} \text{erg s}^{-1} \text{cm}^{-2}$, which might not be detectable. A tantalizing possibility that the disk dominates the emission starting as early as 10^4 s, at which point the source flux equals the Eddington flux of a hypothesized IMBH, is unrealistic. First, the optical depth of supernova debris is about $\tau \sim 10^4$ at that time, and the special directions of low $\tau \lesssim 1$ are not expected. Second, the accretion rate at 10^4 s is expected to be highly super-Eddington, if only a tiny fraction of the WD material is bound to the BH. Third, even if the accretion rate at 10^4 s was at the Eddington level, then the extrapolated jet power at 10^3 s for the standard jet efficiency would be ~ 100 times lower than the true soft X-ray power for a jet with $\theta = 0.5$ opening angle.

5. EVENT RATES

The IMBHs are thought to reside either in the GCs or in the dwarf galaxies. As we discussed in the earlier work (Haas et al. 2012), the disruption rate of stars by IMBHs in GCs is very uncertain. Baumgardt et al. (2004) predicts that for a GC with a $10^3 M_{\odot}$ central IMBH, the optimistic disruption rate of stars is 10^{-7}yr^{-1} per GC, while 15% of all disruptions are those of WDs. Then, following McLaughlin (1999) we theoretically estimate the space density of the GCs to be

$$n_{\text{GC,th}} \approx 34 \text{Mpc}^{-3}, \quad (18)$$

which leads to the rate (Haas et al. 2012) $R_{\text{WD-IMBH}} \sim 500 \text{yr}^{-1} \text{Gpc}^{-3}$ of the WD disruptions by the IMBHs. Observational constraints on the GC population suggest a slightly lower space density of the GCs (Brodie & Strader 2006; Ramirez-Ruiz & Rosswog 2009)

$$n_{\text{GC,obs}} \approx 4 \text{Mpc}^{-3}, \quad (19)$$

so that the rate estimate should be revised down to

$$R_{\text{WD-IMBH}} \sim 50 \text{yr}^{-1} \text{Gpc}^{-3}. \quad (20)$$

The disruption rate of stars in dwarf galaxies was estimated to be very high. If $M - \sigma$ (Gebhardt et al. 2000) and $M_{\text{BH}} - M_{\text{bulge}}$ (Magorrian et al. 1998; Marconi & Hunt 2003) relations hold down to the low BH masses, then the predicted tidal disruption rate is higher for the dwarf galaxies compared to the other galaxies (Wang & Merritt 2004). Dependent on the radial stellar profile the disruption rate in a dwarf galaxy with a $10^4 M_{\odot}$ BH varies from $R_{\text{IMBH}} = 10^{-5} \text{yr}^{-1}$ per galaxy up to $R_{\text{IMBH}} = 10^{-1} \text{yr}^{-1}$ per galaxy. For an estimate we take the rate of $R_{\text{IMBH}} = 10^{-3} \text{yr}^{-1}$ per galaxy. Note that during a tidal disruption some fraction f of a WD accretes onto a BH. Then the mass of an IMBH grows at a rate

$$\frac{dM_{\text{BH}}}{dt} = f R_{\text{IMBH}} m_{\text{WD}}. \quad (21)$$

Taking a large fiducial fraction $f = 0.5$ we find that the BH mass doubles every $t_d = 2 \times 10^8$ years solely due to tidal disruptions of WDs, where we assumed that the same $R_{\text{WD-IMBH}}/R_{\text{IMBH}} = 15\%$ of the disrupted stars are the WDs. This doubling is consistent with the estimated age of GRB060218 host galaxy 200 Myr (Ferrero et al. 2007). As we will show in the next subsection, an IMBH swallows only a small fraction of a MS star, so that the IMBH mass growth rate due to the MS stars may not be dominant.

According to observations (Ferguson & Sandage 1991) the number density of dwarf d galaxies with V magnitude $M_V > -15.5$ is about the number density of large elliptical E and spiral S galaxies. Thus, the space density of dwarf galaxies is $n_{\text{dwarf}} \sim 10^{-2} \text{Mpc}^{-3}$. Then, assuming that every dwarf galaxy has a small enough IMBH, we arrive at a rate estimate

$$R_{\text{WD-IMBH}} \sim 1500 \text{yr}^{-1} \text{Gpc}^{-3}, \quad (22)$$

which is much larger than the rate in GCs. One event per year for the field of view (FOV) 10% of the sky would happen as close as 200Mpc . Such disruptions should be detected by *Swift* BAT instrument with FOV of approximately 10% of the sky. Collimation into a wide jetted outflow may reduce the rate down to one event per several years consistently with a single GRB060218 source at 143 Mpc distance.

Since the WD disruptions by the IMBHs produce supernovae, we can compare the rate of those with the total supernova rates. Li et al. (2011) find the local rate of supernovae Type Ia to be

$$R_{\text{SN,Ia}} \approx 3 \times 10^4 \text{yr}^{-1} \text{Gpc}^{-3}, \quad (23)$$

and the rate of supernovae Type Ib/c to be $R_{\text{SN,Ib/c}} \approx 3 \times 10^4 \text{yr}^{-1} \text{Gpc}^{-3}$. In an optimistic estimate, if a supernova is produced in every tidal disruption of a WD, then one disruption-induced supernova happens per 20 Type Ia supernovae. A distinctive feature of disruption-induced supernovae is their location in the galactic nuclei substantially close to the photometric centers of their dwarf host galaxies. This rate equals to the rate of calcium-rich gap transients, whose properties may be explained by the underluminous explosions of the WDs (Kasliwal et al. 2012). However, those transients occur very far from the galactic nuclei.

5.1. Disruptions of MS Stars by IMBHs

In the present paper we consider the radiative signatures of the WD disruptions by the IMBHs. However, the disruptions rates of the MS stars by the IMBHs are several times higher. Then the question is whether these more frequent events produce equally distinct and observable signatures. In this subsection we show that MS star/IMBH disruptions are quite different: they are much longer and much fainter. Despite lower rates, it might be easier to observe the WD/IMBH disruptions compared to the MS star/IMBH disruptions.

The fallback timescale is about $t_{\text{fb}} \sim 20$ days for a disruption of a $1 M_{\odot}$ MS star with a solar radius by a $10^4 M_{\odot}$ IMBH with a pericenter radius equal to the tidal radius $R_p = R_T$. The peak fallback rate is $\dot{M}_{\text{peak}} \sim 5 M_{\odot} \text{yr}^{-1}$. While a tidal disruption radius is $R_T \sim 20 R_g$ for a WD, a MS star with the solar mass and the solar radius has a much larger $R_T \sim 1000 R_g$. Since the peak fallback accretion rate $\dot{M}_{\text{peak}} \sim 5 M_{\odot} \text{yr}^{-1}$ is much above the Eddington rate $\dot{M}_{\text{Edd}} \sim 2 \times 10^{-4} M_{\odot} \text{yr}^{-1}$, the RIAF settles from the outer disk radius $R_{\text{disk}} \approx 2 R_p$ down to the BH. A RIAF with an outer radius greater than about $20 R_g$ settles into the convection dominated accretion flow (CDAF) (Narayan et al. 2000; Quataert & Gruzinov 2000), where the energy transport between the inner and the outer flow starts to play a role. Unlike the fallback disks following the WD disruptions, the disks following the disruptions of the MS stars always exist in CDAF state. A flow with convection has a shallow density profile $\rho \propto r^{-\beta}$ down to the inner radius $R_{\text{in}} \sim 20 R_g$ (Abramowicz et al. 2002). Only a small fraction

of available matter accretes in CDAF state. The sustained BH accretion rate is

$$\dot{M} = \dot{M}_{\text{fb}} (R_{\text{disk}}/R_{\text{in}})^{1.5-\beta}. \quad (24)$$

The density slope $\beta = 1.5 - s = 0.5 - 1.0$ was found in the numerical simulations (see Yuan et al. 2012 for the review). The value $\beta = 0.8 - 0.9$ was estimated by Shcherbakov et al. (2012) for Sgr A*. The flow settles on a viscous timescale $t_{\text{visc}} \sim 1$ day, which is shorter than $t_{\text{fb}} \sim 20$ days. Thus, the peak accretion rate of the CDAF is given by the formula (24) with $\dot{M}_{\text{fb}} = \dot{M}_{\text{fb,peak}}$. The peak accretion rate onto the BH can be a factor of 10–100 lower than the peak fallback rate

$$\dot{M}_{\text{peak}} \sim (0.01 - 0.1) \dot{M}_{\text{fb,peak}}. \quad (25)$$

Let us make a fiducial estimate for a BH mass $10^4 M_{\odot}$ and a disruption of a Sun-like star for $R_p = R_T$ assuming density slope $\beta = 0.85$. In this case the accretion rate is lowered by a factor of ≈ 23 due to the action of convection according to the formula (24), and the resultant jet power is $L_{\text{kin}} = 0.01 \dot{M}_{\text{peak}} c^2 \sim 10^{44} \text{erg s}^{-1}$ with the same fiducial efficiency $P_0 \sim 0.01$ used in estimates for the WD disruptions. A jet with such low power has the photospheric radius smaller than the saturation radius. The radiation decouples from matter before the jet can accelerate, and the powerful photospheric emission is not expected. The photospheric radius corresponds to the bulk Lorentz factor $\Gamma \approx 2$. If the source is beamed towards us, then the observed temperature is $T_{\text{ob}} \approx 0.5$ keV. The radiated blackbody luminosity is $L_{\text{BB}} \sim 1 \times 10^{43} \text{erg s}^{-1}$, which corresponds to the distance $d \sim 10$ Mpc to match the peak flux level $F_{\text{peak}} \approx 10^{-8} \text{erg s}^{-1} \text{cm}^{-2}$ of GRB060218. The absorption by the cooled down jet exhaust or the dense wind, which consist mostly of hydrogen and helium, may substantially lower the X-ray luminosity and turn the source into an optical/UV transient similar to the disruption of a MS star by a SMBH (Strubbe & Quataert 2009; Strubbe & Quataert 2011). The optical/UV transients lasting for 20 days and releasing up to $10^{49} \text{erg s}^{-1}$ may be confused with variability of the AGNs. According to the standard theory (Cannizzo & Gehrels 2009; Cannizzo et al. 2011) at $t_x \sim 100$ days the source should switch from the fallback-dominated $t^{-5/3}$ behavior to the disk spreading behavior $t^{-4(1+s)/3} \propto t^{-2.2}$. Yet, the timescale t_x was derived for the adiabatic RIAFs. The temporal behavior of radiatively inefficient fallback accretion disks with convection were not explored to our knowledge and should be investigated in the future work. If absorption is inefficient, then one transient per year is expected at a distance $d \sim 100$ Mpc with the peak flux level of about $F_{\text{peak}} \approx 10^{-10} \text{erg s}^{-1} \text{cm}^{-2}$. Searches within the existing and future X-ray and optical surveys should identify such sources and constrain the rates of tidal disruptions of various kinds. ROSAT All-sky survey search revealed several nearby tidal disruption candidates with redshifts $z \lesssim 0.1$ (Komossa 2002). Some of those candidates, e.g. RX J1242-1119, cannot be easily matched with a large host galaxy (Komossa & Greiner 1999; Komossa et al. 2004). The absence of suitable X-ray/optical transients in the data might not immediately invalidate high tidal disruption rates. For example, if the jet power in the disruptions of the MS stars is as weak as in GRB060218, then the X-ray/optical flux from such disruptions is expected to be up to 100 times lower. The detailed analysis of the tidal disruptions of the MS stars is beyond the scope of the present manuscript.

A candidate source for a stellar disruption by an IMBH is a

source in a GC NGC1399 explored by Irwin et al. (2010). Observed at a distance 20 Mpc the source exhibited X-ray luminosity of $2 \times 10^{39} \text{ erg s}^{-1}$ and luminosity of a few $\times 10^{36} \text{ erg s}^{-1}$ in oxygen and nitrogen optical lines. These properties were found consistent with a disruption of a horizontal branch star by a $50\text{--}100M_{\odot}$ BH (Clausen et al. 2012).

6. DISCUSSION AND CONCLUSIONS

In this paper we explore the theory of the WD tidal disruptions by the IMBHs, an understudied, but a very promising type of an encounter. The high rate of the WD tidal disruptions in dwarf galaxies warrants the search for candidates among already observed objects. We identify GRB060218 with the accompanying supernova SN2006aj as a promising candidate and model its temporal/spectral properties and the properties of the host galaxy. GRB060218 was explored previously within the supernova shock breakout model and the model, where a jet is launched by a newborn BH or a neutron star. We find the tidal disruption model to be a viable alternative, which more naturally explains some features of the candidate. In this section we briefly discuss the application of the tidal disruption model to GRB060218/SN2006aj source, and compare different models of the source.

The tidal disruption model performs equally well compared to other models in terms of event rates. The appearance of GRB060218 at a redshift $z = 0.033$ is consistent with a disruption rate 10^{-3} yr^{-1} of all stars in a dwarf galaxy. The rate of low-luminosity GRBs such as our source was estimated to be about 10% of supernova Type Ib/c rate (Toma et al. 2007). A number of shock breakouts were observed to accompany supernovae at all distances (Woosley et al. 1999; Mazzali et al. 2006a; Chevalier & Fransson 2008; Schawinski et al. 2008), and the sample of shock breakouts is likely incomplete.

The long emission duration 2600 s provides a mass estimate $1 \times 10^4 M_{\odot}$ for an IMBH within a tidal disruption scenario. The relativistic shock breakout model also explains the event duration (Nakar & Sari 2012). The extreme smoothness of the spectrum is consistent with either the supernova shock breakout or the jet launched by an IMBH. The jet launched by a stellar mass central object may produce more time variability.

While the observed absorbed X-ray lightcurve of the event can be fitted with the exponential decay within a shock breakout model (Campana et al. 2006), the source flux corrected for the absorption exhibits $(t - t_0)^{-5/3}$ behavior specific to the tidal disruptions. Moreover, the full soft X-ray lightcurve including the rise and the decay phases can be fitted well with the accretion rate temporal dependence $\dot{M}(t)$ for a tidal disruption. The steep flux decay at ~ 6500 s can be ascribed to the source being more heavily absorbed at late times, while the jet power and the radiation power still scale as $t^{-5/3}$. The power-law decay is not expected in a shock breakout, while a shallower temporal slope is expected for a jet-powered source (Soderberg et al. 2006). The temporal slope of the afterglow phase, powered by the central engine, is consistent with the accretion rate temporal dependence $t^{-4/3}$ in a RIAF dominated by the disk evolution behavior.

The observed blackbody flux and the blackbody temper-

ature can be readily modeled by the jet photospheric emission to provide the estimates for the IMBH mass $1 \times 10^4 M_{\odot}$ and the Lorentz factor $\Gamma \sim 3$. Moreover, these estimates stay consistent between different time slices in our time-resolved spectroscopic analysis. The rise of the blackbody temperature accompanied by the constant blackbody flux and the dropping total flux leads to the same BH mass and Lorentz factor over many time slices within the fireball model. The relativistic shock in the wind region of a WR star is able to reproduce the spectrum (Waxman et al. 2007). However, the properties of the underlying WR star and its wind would be peculiar (Li 2007).

The absence of the jet break is consistent with the wide outflow produced without the pressure support from the surrounding star. Both the supernova shock breakout and the tidal disruption scenario produce wide outflows.

Strong early optical/UV emission is consistent with coming from the jet front. The non-relativistic outer shells producing early optical/UV blackbody emission violate the source energetics (Ghisellini et al. 2007b) within the supernova shock breakout model. In addition, the total radiated energy $10^{49.5} \text{ erg}$ of GRB060218 is somewhat higher than expected for the shock breakouts (Colgate 1974) in general.

The host of GRB060218 is a dwarf galaxy, and the dwarf galaxies are expected to have the highest tidal disruption rates. Only the dwarf galaxies and not the bigger galaxies are expected to have the IMBHs in their nuclei, and thus disrupt the WDs. We crudely estimate the central BH mass $1 \times 10^4 M_{\odot}$ based on the stellar mass $10^{7.2} M_{\odot}$ of the host. The supernova position is consistent with the center of the host dwarf galaxy, as it should be for the tidal disruption scenario. Supernovae unrelated to the central BHs generally have large displacements from the galactic centers. The supernova SN2006aj has the low estimated ejecta mass $M_{\text{ej}} = 1 - 2 M_{\odot}$ consistent with the heavy WD.

In sum, the tidal disruption model can explain all features of the source and only the scale of the jet power remains somewhat arbitrary. Note that the BH mass $1 \times 10^4 M_{\odot}$ is estimated in three entirely independent ways! A conclusion that GRB060218 was a tidal disruption by the IMBH in a nucleus of a dwarf galaxy at a distance 150 Mpc indicates that the IMBHs are abundant in the local Universe.

7. ACKNOWLEDGEMENTS

The authors are thankful to Nathaniel Bulter, John Cannizzo, Coleman Miller, Ramesh Narayan, Eve Ostriker, Enrico Ramirez-Ruiz, and Leslie Sage for useful discussions and comments, Special thanks to Abderahmen Zoghbi for advice on temporal analysis, Ranjan Vasudevan for general help with XSPEC, Demosthenes Kazanas and Jonathan McKinney for discussions of the large-scale magnetic field generation. RVS is supported by NASA Hubble Fellowship grant HST-HF-51298.01, PL is supported by NSF grants 1212433, 1205864, 0941417, 0903973, 0855423, and 0903973, CSR thanks NASA for support under Astrophysics Theory Program (ATP) grant NNX10AE41G.

REFERENCES

- Abramowicz, M. A., Czerny, B., Lasota, J. P., & Szuszkiewicz, E. 1988, *ApJ*, 332, 646
 Abramowicz, M. A., Igmenshchev, I. V., Quataert, E., & Narayan, R. 2002, *ApJ*, 565, 1101
 Angel, J. R. P. 1978, *Ann. Rev. Astron. Astr.*, 16, 487
 Arnaud, K. A. 1996, in *Astronomical Society of the Pacific Conference Series*, Vol. 101, *Astronomical Data Analysis Software and Systems V*, ed. G. H. Jacoby & J. Barnes, 17

- Balbus, S. A., & Hawley, J. F. 1991, *ApJ*, 376, 214
 —. 1998, *Reviews of Modern Physics*, 70, 1
- Bardeen, J. M., Press, W. H., & Teukolsky, S. A. 1972, *ApJ*, 178, 347
- Baumgardt, H., Makino, J., & Ebisuzaki, T. 2004, *ApJ*, 613, 1143
- Begelman, M. C., Volonteri, M., & Rees, M. J. 2006, *MNRAS*, 370, 289
- Beloborodov, A. M. 2011, *ApJ*, 737, 68
- Blandford, R. D., & Payne, D. G. 1982, *MNRAS*, 199, 883
- Blandford, R. D., & Znajek, R. L. 1977, *MNRAS*, 179, 433
- Bloom, J. S., Kulkarni, S. R., & Djorgovski, S. G. 2002, *AJ*, 123, 1111
- Bloom, J. S., et al. 2011, *Science*, 333, 203
- Brandenburg, A., & Subramanian, K. 2005, *Physics Reports*, 417, 1
- Brodie, J. P., & Strader, J. 2006, *Annual Rev. of A. and A.*, 44, 193
- Bromberg, O., Nakar, E., & Piran, T. 2011, *ApJ*, 739, L55
- Bufano, F., et al. 2012, *ApJ*, 753, 67
- Burrows, D. N., et al. 2011, *Nature*, 476, 421
- Butler, N. R. 2007, *ApJ*, 656, 1001
- Butler, N. R., Kocevski, D., Bloom, J. S., & Curtis, J. L. 2007, *ApJ*, 671, 656
- Campana, S., et al. 2006, *Nature*, 442, 1008
- Cannizzo, J. K., & Gehrels, N. 2009, *ApJ*, 700, 1047
- Cannizzo, J. K., Lee, H. M., & Goodman, J. 1990, *ApJ*, 351, 38
- Cannizzo, J. K., Troja, E., & Lodato, G. 2011, *ApJ*, 742, 32
- Chevalier, R. A., & Fransson, C. 2008, *ApJ*, 683, L135
- Clausen, D., Sigurdsson, S., Eracleous, M., & Irwin, J. A. 2012, *MNRAS*, 424, 1268
- Colbert, E. J. M., & Mushotzky, R. F. 1999, *ApJ*, 519, 89
- Colgate, S. A. 1968, *Canadian Journal of Physics*, 46, 476
 —. 1974, *ApJ*, 187, 333
- Daigne, F., & Mochkovitch, R. 2002, *MNRAS*, 336, 1271
- Dainotti, M. G., Bernardini, M. G., Bianco, C. L., Caito, L., Guida, R., & Ruffini, R. 2007, *A&A*, 471, L29
- Daly, R. A. 2011, *MNRAS*, 414, 1253
- Davis, S. W., Narayan, R., Zhu, Y., Barret, D., Farrell, S. A., Godet, O., Servillat, M., & Webb, N. A. 2011, *ApJ*, 734, 111
- Davis, S. W., Stone, J. M., & Pessah, M. E. 2010, *ApJ*, 713, 52
- De Colle, F., Guillochon, J., Naiman, J., & Ramirez-Ruiz, E. 2012, *ApJ*, 760, 103
- Dong, X., et al. 2007, *ApJ*, 657, 700
- Eichler, D., & Levinson, A. 2000, *ApJ*, 529, 146
- Evans, C. R., & Kochanek, C. S. 1989, *ApJ*, 346, L13
- Fabbiano, G., Schweizer, F., & Mackie, G. 1997, *ApJ*, 478, 542
- Fabbiano, G., Zezas, A., & Murray, S. S. 2001, *ApJ*, 554, 1035
- Fan, Y.-Z., Piran, T., & Xu, D. 2006, *J. of Cosmology and Astroparticle Phys.*, 9, 13
- Ferguson, H. C., & Sandage, A. 1991, *AJ*, 101, 765
- Ferrero, P., Palazzi, E., Pian, E., & Savaglio, S. 2007, in *American Institute of Physics Conference Series*, Vol. 924, *The Multicolored Landscape of Compact Objects and Their Explosive Origins*, ed. T. di Salvo, G. L. Israel, L. Piersant, L. Burderi, G. Matt, A. Tornambe, & M. T. Menna, 120–125
- Ferrero, P., et al. 2006, *A&A*, 457, 857
- Fryer, C. L., Woosley, S. E., & Heger, A. 2001, *ApJ*, 550, 372
- Gebhardt, K., Rich, R. M., & Ho, L. C. 2002, *ApJ*, 578, L41
- Gebhardt, K., et al. 2000, *ApJ*, 539, L13
- Ghisellini, G. 2012, in *-Ray Bursts 2012 Conference (GRB 2012)*
- Ghisellini, G., Ghirlanda, G., & Tavecchio, F. 2007a, *MNRAS*, 382, L77
 —. 2007b, *MNRAS*, 375, L36
- Giannios, D. 2006, *A&A*, 457, 763
 —. 2012, *MNRAS*, 422, 3092
- Gorosabel, J., et al. 2006, *A&A*, 459, L33
- Greene, J. E. 2012, *ArXiv e-prints*
- Guillochon, J., & Ramirez-Ruiz, E. 2012, *arXiv:1206.2350*
- Gültekin, K., Miller, M. C., & Hamilton, D. P. 2004, *ApJ*, 616, 221
- Haas, R., Shcherbakov, R. V., Bode, T., & Laguna, P. 2012, *ApJ*, 749, 117
- Irwin, J. A., Brink, T. G., Bregman, J. N., & Roberts, T. P. 2010, *ApJ*, 712, L1
- Jorstad, S. G., et al. 2005, *AJ*, 130, 1418
- Kasliwal, M. M., et al. 2012, *ApJ*, 755, 161
- Katz, B., Budnik, R., & Waxman, E. 2010, *ApJ*, 716, 781
- Khokhlov, A., Mueller, E., & Hoefflich, P. 1993, *A&A*, 270, 223
- King, A. R., Pringle, J. E., & Livio, M. 2007, *MNRAS*, 376, 1740
- Komossa, S. 2002, in *Lighthouses of the Universe: The Most Luminous Celestial Objects and Their Use for Cosmology*, ed. M. Gilfanov, R. Sunyaev, & E. Churazov, 436
- Komossa, S., & Greiner, J. 1999, *A&A*, 349, L45
- Komossa, S., Halpern, J., Schartel, N., Hasinger, G., Santos-Lleo, M., & Predehl, P. 2004, *ApJ*, 603, L17
- Laguna, P., Miller, W. A., Zurek, W. H., & Davies, M. B. 1993, *ApJ*, 410, L83
- Levan, A. J., et al. 2011, *Science*, 333, 199
- Li, L.-X. 2007, *MNRAS*, 375, 240
- Li, W., Chornock, R., Leaman, J., Filippenko, A. V., Poznanski, D., Wang, X., Ganeshalingam, M., & Mannucci, F. 2011, *MNRAS*, 412, 1473
- Liang, E.-W., Yi, S.-X., Zhang, J., Lü, H.-J., Zhang, B.-B., & Zhang, B. 2010, *ApJ*, 725, 2209
- Liu, D., Pe'er, A., & Loeb, A. 2012, *arXiv:1211.5120*
- Lodato, G., King, A. R., & Pringle, J. E. 2009, *MNRAS*, 392, 332
- Luminet, J.-P. 1985, *Annales de Physique*, 10, 101
- Luminet, J.-P., & Marck, J.-A. 1985, *MNRAS*, 212, 57
- Luminet, J.-P., & Pichon, B. 1989a, *A&A*, 209, 103
 —. 1989b, *A&A*, 209, 85
- Madau, P., & Rees, M. J. 2001, *ApJ*, 551, L27
- Maeda, K., et al. 2007, *ApJ*, 658, L5
- Magorrian, J., et al. 1998, *AJ*, 115, 2285
- Marconi, A., & Hunt, L. K. 2003, *ApJ*, 589, L21
- Matsumoto, H., Tsuru, T. G., Koyama, K., Awaki, H., Canizares, C. R., Kawai, N., Matsushita, S., & Kawabe, R. 2001, *ApJ*, 547, L25
- Mazzali, P. A., et al. 2006a, *ApJ*, 645, 1323
 —. 2006b, *Nature*, 442, 1018
 —. 2007, *ApJ*, 661, 892
 —. 2008, *Science*, 321, 1185
- McKinney, J. C. 2005, *ApJ*, 630, 5
- McKinney, J. C., & Blandford, R. D. 2009, *MNRAS*, 394, L126
- McKinney, J. C., Tchekhovskoy, A., & Blandford, R. D. 2012, *MNRAS*, 423, 3083
- McLaughlin, D. E. 1999, *AJ*, 117, 2398
- McMillan, P. J. 2011, *MNRAS*, 414, 2446
- Meszáros, P. 2006, *Rep. on Progress in Phys.*, 69, 2259
- Mészáros, P., & Rees, M. J. 2000, *ApJ*, 530, 292
- Metzger, B. D. 2012, *MNRAS*, 419, 827
- Metzger, B. D., Giannios, D., & Mimica, P. 2012, *MNRAS*, 420, 3528
- Miller-Jones, J. C. A., Fender, R. P., & Nakar, E. 2006, *MNRAS*, 367, 1432
- Mirabal, N., Halpern, J. P., An, D., Thorstensen, J. R., & Terndrup, D. M. 2006, *ApJ*, 643, L99
- Misra, K., Fruchter, A. S., & Nugent, P. 2011, in *American Institute of Physics Conference Series*, Vol. 1358, *American Institute of Physics Conference Series*, ed. J. E. McEnery, J. L. Racusin, & N. Gehrels, 299–302
- Modjaz, M., et al. 2006, *ApJ*, 645, L21
- Nakar, E., & Sari, R. 2012, *ApJ*, 747, 88
- Należyty, M., & Madej, J. 2004, *A&A*, 420, 507
- Narayan, R., Igmenshchev, I. V., & Abramowicz, M. A. 2000, *ApJ*, 539, 798
- Narayan, R., Piran, T., & Kumar, P. 2001, *ApJ*, 557, 949
- Nauenberg, M. 1972, *ApJ*, 175, 417
- Nava, L., Ghirlanda, G., Ghisellini, G., & Firmani, C. 2008, *MNRAS*, 391, 639
- Nishimura, J., Mitsuda, K., & Itoh, M. 1986, *PASJ*, 38, 819
- O'Neill, S. M., Reynolds, C. S., Miller, M. C., & Sorathia, K. A. 2011, *ApJ*, 736, 107
- Papadakis, I. E., & Lawrence, A. 1993, *MNRAS*, 261, 612
- Pe'er, A. 2008, *ApJ*, 682, 463
- Pe'er, A., Mészáros, P., & Rees, M. J. 2005, *ApJ*, 635, 476
 —. 2006, *ApJ*, 642, 995
- Pe'er, A., & Ryde, F. 2011, *ApJ*, 732, 49
- Pe'er, A., Ryde, F., Wijers, R. A. M. J., Mészáros, P., & Rees, M. J. 2007, *ApJ*, 664, L1
- Penna, R. F., McKinney, J. C., Narayan, R., Tchekhovskoy, A., Shafee, R., & McClintock, J. E. 2010, *MNRAS*, 408, 752
- Pian, E., et al. 2006, *Nature*, 442, 1011
- Poutanen, J., & Svensson, R. 1996, *ApJ*, 470, 249
- Pushkarev, A. B., Kovalev, Y. Y., Lister, M. L., & Savolainen, T. 2009, *A&A*, 507, L33
- Putney, A. 1999, in *Astronomical Society of the Pacific Conference Series*, Vol. 169, *11th European Workshop on White Dwarfs*, ed. S.-E. S. E. G. Meistas, 195
- Quataert, E., & Gruzinov, A. 2000, *ApJ*, 539, 809
- Quataert, E., & Kasen, D. 2012, *MNRAS*, 419, L1
- Ramirez-Ruiz, E., & Rosswog, S. 2009, *ApJ*, 697, L77
- Rees, M. J. 1984, *Annual Rev. of Astronomy and Astrophys.*, 22, 471
- Rees, M. J., & Meszaros, P. 1994, *ApJ*, 430, L93
- Rees, M. J., & Mészáros, P. 2005, *ApJ*, 628, 847
- Rosswog, S., Ramirez-Ruiz, E., & Hix, R. 2009, *ApJ*, 695, 404

- Rybicki, G. B., & Lightman, A. P. 1979, *Radiative Processes in Astrophysics* (New York: Wiley)
- Sari, R., & Piran, T. 1997, *ApJ*, 485, 270
- Schawinski, K., et al. 2008, *Science*, 321, 223
- Schneider, R., Ferrara, A., Natarajan, P., & Omukai, K. 2002, *ApJ*, 571, 30
- Schramm, M., & Silverman, J. D. 2012, arXiv:1212.2999
- Shakura, N. I., & Sunyaev, R. A. 1973, *Astron. and Astrophys.*, 24, 337
- Shapiro, S. L., & Teukolsky, S. A. 1986, *Black Holes, White Dwarfs and Neutron Stars: The Physics of Compact Objects* (Wiley: New York)
- Shcherbakov, R. V., Penna, R. F., & McKinney, J. C. 2012, *ApJ*, 755, 133
- Soderberg, A. M., et al. 2006, *Nature*, 442, 1014
- . 2008, *Nature*, 453, 469
- Stone, J. M., Hawley, J. F., Gammie, C. F., & Balbus, S. A. 1996, *ApJ*, 463, 656
- Strubbe, L. E., & Quataert, E. 2009, *MNRAS*, 400, 2070
- Strubbe, L. E., & Quataert, E. 2011, *MNRAS*, 415, 168
- Svensson, K. M., Levan, A. J., Tanvir, N. R., Fruchter, A. S., & Strolger, L.-G. 2010, *MNRAS*, 405, 57
- Takiwaki, T., Kotake, K., & Sato, K. 2009, *ApJ*, 691, 1360
- Tchekhovskoy, A., Narayan, R., & McKinney, J. C. 2010, *ApJ*, 711, 50
- Tchekhovskoy, A., Narayan, R., & McKinney, J. C. 2010, *New Astronomy*, 15, 749
- . 2011, *MNRAS*, 418, L79
- Titarchuk, L. 1994, *ApJ*, 434, 313
- Toma, K., Ioka, K., Sakamoto, T., & Nakamura, T. 2007, *ApJ*, 659, 1420
- Usov, V. V. 1992, *Nature*, 357, 472
- Uzdensky, D. A., & MacFadyen, A. I. 2007, *ApJ*, 669, 546
- Wang, J., & Merritt, D. 2004, *ApJ*, 600, 149
- Wang, X., et al. 2009, *ApJ*, 699, L139
- Waxman, E., Mészáros, P., & Campana, S. 2007, *ApJ*, 667, 351
- Wheeler, J. C., Yi, I., Höflich, P., & Wang, L. 2000, *ApJ*, 537, 810
- Wiersema, K., et al. 2007, *A&A*, 464, 529
- Woosley, S. E., Eastman, R. G., & Schmidt, B. P. 1999, *ApJ*, 516, 788
- Yuan, F., Wu, M., & Bu, D. 2012, *ApJ*, 761, 129
- Zauderer, B. A., et al. 2011, *Nature*, 476, 425